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Six months of mass outflow and inclined rings in the ejecta of V1494 Aql

S. P. S. Eyres,^{1★} I. Heywood,² T. J. O’Brien,² R. J. Ivison,³ T. W. B. Muxlow²
and V. G. Elkin¹

¹*Centre for Astrophysics, University of Central Lancashire, Preston PR1 2HE*

²*Jodrell Bank Observatory, University of Manchester, Macclesfield, Cheshire SK11 9DL*

³*Astronomy Technology Centre, Royal Observatory, Blackford Hill, Edinburgh EH9 3HJ*

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ABSTRACT

V1494 Aql was a very fast nova which reached a visual maximum of $m_v \simeq 4.0$ by the end of 1999 December 3. We report observations from 4 to 284 d after discovery, including submillimetre- and centimetre-band fluxes, a single MERLIN image and optical spectroscopy in the 410 to 700 nm range. The extent of the radio continuum emission is consistent with a recent lower distance estimate of 1.6 kpc. We conclude that the optical and radio emission arises from the same expanding ejecta. We show that these observations are not consistent with simple kinematical spherical shell models used in the past to explain the rise and fall of the radio flux density in these objects. The resolved remnant structure is consistent with an inclined ring of enhanced density within the ejecta. Optical spectroscopy indicates likely continued mass ejection for over 195 d, with the material becoming optically thin in the visible sometime between 195 and 285 d after outburst.

Key words: stars: individual: V1494 Aql – novae, cataclysmic variables – stars: winds, outflows – radio continuum: stars.

1 INTRODUCTION

When a classical nova undergoes an eruption, the initial visual brightening is accompanied by the ejection of perhaps 10^{-5} to $10^{-4} M_\odot$ of matter. This expands away from the central binary, and as the visual light starts to decline the increasing solid angle of the ionized ejecta leads to steadily rising, optically thick, thermal radio emission. As the remnant expands the emission becomes optically thin at ever longer wavelengths, and fades below detectability. This process is understood in general terms, but the development of structure in the remnant observed a few months after outburst is still unclear. Asymmetric ejection, intra-ejecta interaction and the influence of the binary companion all have a role to play.

Much still needs to be done observationally to constrain these models. Examining the very early development of the remnant is one means by which this is currently taking place. Radio observations hold out the possibility of imaging the remnant structure at very early times, particularly with MERLIN providing 50-mas resolution. Observations to date (Eyres, Davis & Bode 1996; Eyres et al. 2000) demonstrate that the clumpy, non-spheroidal structures seen at late times (e.g. Gill & O’Brien 1998, 2000) appear very early, and also develop rapidly. For V1974 Cyg, in which the shell structure was resolved by MERLIN only 80 days after outburst (Pavelin et al. 1993; Eyres et al. 1996), Ivison et al. (1993) found that the

submillimetre to millimetre data early in the outburst were inconsistent with kinematical models used to fit radio data at later times. In many novae optical spectroscopy is consistent with axisymmetric structures, with features due to density enhancements in polar caps, tropical rings and equatorial bands evident to different degrees in different novae (e.g. Gill & O’Brien 1999). In V705 Cas the radio images show the structures developing asymmetrically while the optical spectra are consistent with an axisymmetric density distribution (Eyres et al. 2000). V723 Cas showed a similar development to V1974 Cyg (Heywood et al. 2005). Observing further novae over a range of speed classes [see Payne-Gaposchkin (1964) for a definition of these classes] and hence ejecta velocities is important in constraining revised models over the physical parameter space.

V1494 Aql was first detected at $m_v \simeq 6.0$ on 1999 December 1 (Pereira et al. 1999), reaching visual maximum of $m_v \simeq 4.0$ by the end of 1999 December 3 (Rao et al. 1999; Liller, Frasca & Messina 1999). It dropped to fainter than $m_v \simeq 6.0$ by 1999 December 10 (Pontefract et al. 1999) and $m_v \simeq 7.0$ by 1999 December 19 (Keen et al. 1999). Kiss & Thomson (2000) find $t_2 = 6.6 \pm 0.5$ d and $t_3 = 16 \pm 0.5$ d, making it a very fast nova. Optical spectroscopy has identified velocity components at up to 1700 km s^{-1} (Fujii, Ayani & Liller 1999), while infrared spectroscopy shows velocities up to 2900 km s^{-1} (Rudy et al. 2001), consistent with this classification. Iijima & Esenoglu (2003) concluded that mass ejection continued for over 150 d, contradicting a more general conclusion that ejection in novae ceases after a few days (e.g. Hjellming 1996). In this work we present the results of observations of V1494 Aql starting on 1999

★E-mail: spseyres@uclan.ac.uk

December 7, within 5 d of its discovery in outburst, and interpret them in the context of standard models of millimetre–centimetre radio emission from novae.

2 OBSERVATIONS

We report observations in the radio, submillimetre and optical bands using the Multi-Element Radio-Linked Interferometer Network (MERLIN), the Very Large Array (VLA), the James Clerk Maxwell Telescope (JCMT) and the 1-m Zeiss telescope at the Special Astrophysical Observatory (SAO) of the Russian Academy of Science. Observing dates range from 4 to 288 d after visual maximum (1999 December 3, taken as day 0 hereafter).

2.1 JCMT

Initial observations were made with the JCMT on 1999 December 7 and 8 (days 4 and 5), at 450 and 850 μm using the Submillimetre Common-User Bolometer Array (SCUBA) in photometry mode (Ivison et al. 1998; Holland et al. 1999) in good weather (optical depth $\tau \sim 0.07$ at 225 GHz). Calibration was made against the JCMT primary calibrator Uranus, with skydips to measure opacity at both observing wavelengths.

2.2 Radio observations

MERLIN Target-of-Opportunity observations were triggered and took place on 1999 December 30 (day 27) at 18 cm. The observations were made in phase-referenced mode, using 1919 + 063 as the phase-reference source. The flux density scale was set by comparison of the unresolved source OQ 208 with the resolved flux density calibrator 3C 286. On the shortest MERLIN spacings at 1658 MHz, 3C 286 has a flux density of 13.639 Jy. OQ 208, by comparison, has a flux density of 1.125 Jy, giving the calibrated flux density of 1919 + 063 to be 0.164 Jy. The data were mapped in wide-field mode to deal with a bright confusing source in the side-lobes of the primary beams of the antennas. A later MERLIN observation was made on 2000 April 17 (day 136) at 6 cm, using 1929 + 050 as a phase-reference source with a calibrated flux density of 0.27 Jy. The data for V1494 Aql were phase-calibrated using the solutions for the phase-reference source.

Ad hoc VLA data were obtained on 2000 June 7 (day 187) in C-array. Observations were made at 20, 6.1 and 3.5 cm, with a bandwidth of 50 MHz in each case. The primary calibrator was 1331 + 305 (=3C 286), which was used to calibrate the secondary calibrator 1922 + 155, for which flux densities of 0.61 ± 0.04 , 0.632 ± 0.007 and 0.63 ± 0.01 Jy were determined in each band respectively. This was used to flux- and phase-calibrate the observations of V1494 Aql.

2.3 Optical spectroscopy

Observations were made with the Universal Astronomical Grating Spectrograph (UAGS) on the 1-m Zeiss telescope at the SAO on 2000 June 14 and 15 (days 194 and 195) and September 9 and 14 (days 281 and 286). The wavelength calibration was carried out using a neon arc with bespoke software for the reduction of data from this instrument (Vlasyuk 1993). The observations were made with two different gratings, yielding spectral resolutions of around 1000 (June) and 2000 (September), giving velocity resolutions of 300 and 150 km s^{-1} respectively. The data are not flux-calibrated.

These observations were made when V1494 Aql was still bright. We see no evidence for contamination of the spectra by features from the nearby G-type star, as seen at later times by Kiss, Csák & Derekas (2004).

3 RESULTS

On December 7 and 8 (days 4 and 5 after maximum), the flux densities, taking an integration time weighted average over both nights, were 22.5 ± 1.6 mJy at 850 μm and 114 ± 28 mJy at 450 μm , with marginal evidence of an increase between the two nights at 850 μm (a rise of 5.9 mJy compared with a combined uncertainty in the difference of 5.5 mJy). By the time of the MERLIN observations on day 27, the total flux density was 0.50 ± 0.07 mJy at 18 cm. The emission was unresolved at this time, as expected. By day 136 the remnant was partially resolved with MERLIN at 6 cm, and the total detected flux density was found to be 25.0 ± 1.2 mJy (see Fig. 1). The emission detected with MERLIN is consistent with the optical position from Kiss et al. (2004) of $19^{\text{h}}23^{\text{m}}5^{\text{s}}.37$, $+4^{\circ}56'19''.79$ (J2000, quoted uncertainty ± 0.2 arcsec).

On day 187, the VLA detected the nova at 3.5 and 6.1 cm, with peak fluxes and total flux density as given in Table 1. At 20 cm confusing sources make the results difficult to assess, but estimated upper limits on the peak flux and total flux density are also given in Table 1. The emission was unresolved and peak positions in Table 1 were determined from Gaussian fits. These are consistent with both the MERLIN and Kiss et al. (2004) positions, allowing for the relatively low resolution of the VLA in C-array.

Optical spectra were taken on days 194 and 195, and again on days 281 and 286. Fig. 2 shows these spectra, which clearly exhibit emission from H α , H β , H γ , He I, [O III] and [N II]. In addition, the blended lines possibly include He II, [Fe III], [Fe VI], [Fe VII], [Ca V], C IV and Si II, although many of these are tentative identifications owing to the breadth of the features. In all cases where individual

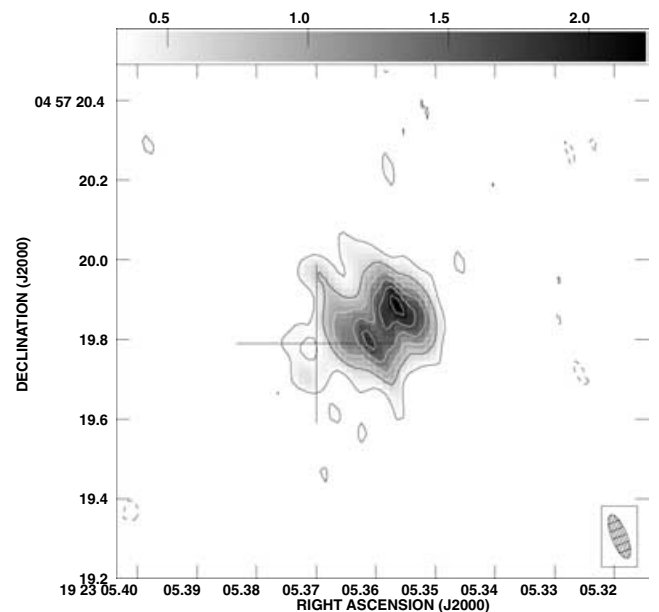


Figure 1. MERLIN image of the remnant of V1494 Aql from 2000 April 17, day 136. The optical position of Kiss et al. (2004) is marked as a cross, with the size indicating the uncertainty. The contours are -3 , 3 , 6 , 9 , 12 , 15 and 18×0.114 mJy beam $^{-1}$ and the grey-scale runs from 0.4 to 2.1 mJy beam $^{-1}$. The peak is 2.10 ± 0.11 mJy beam $^{-1}$.

Table 1. Observed parameters of V1494 Aql on 2000 June 7, day 187, from VLA observations. *Peak positions relative to $19^{\text{h}}23^{\text{m}}00^{\text{s}}, +4^{\circ}57'00''$ (J2000), and consistent with the optical position of Kiss et al. (2004). **These should be regarded as estimated upper limits due to confusing sources.

Wavelength (cm)	Peak (mJy beam ⁻¹)	Total (mJy)	RA* (s)	Dec.* (arcsec)
20	$19 \pm 2^{**}$	$20 \pm 2^{**}$	5.390	20.19
6.1	33.6 ± 0.2	35.4 ± 0.2	5.322	19.71
3.5	56 ± 4	59 ± 4	5.337	19.59

species are reasonably distinct or dominant, the line profiles are consistent with a complex velocity structure in the ejecta. This type of profile has been seen in previous classical novae (Gill & O’Brien 1999), but it is apparent that the width of the lines in V1494 Aql means that adjacent species are blended in most cases. In particular, the $[\text{N II}] \lambda 6583.6\text{-}\text{\AA}$ profile will have a blue peak merged with the $\text{H}\alpha$ profile. Only $\text{H}\beta$ is sufficiently distinct to draw further conclusions. In general the profiles on the two dates show very similar peak wavelengths for each component, symmetric around the rest wavelength. In 2000 June the blueshifted component appears to be suffering absorption with respect to the redshifted component; by 2000 September the two peaks are of comparable strength, consistent with a reduction in the optical depth at this wavelength.

We have carried out five-component Gaussian fits to the $\text{H}\beta$ profiles in order to find the wavelengths of outlying peaks. These values,

along with the estimated full width to half-maximum (FWHM) and full width to zero intensity (FWZI) for each date, are given in Table 2. There is evidence for acceleration in the early decline (Iijima & Esenoglu 2003) but none here for broadening of the profiles with time. The FWHM of the lines are consistent with the values of around 2700 km s^{-1} in 2000 June and September presented by Iijima & Esenoglu (2003). However, there is no evidence for the variable high-velocity components that they found in their spectra of 2000 March.

4 DISCUSSION

4.1 Model fitting

Models for radio emission from novae have generally assumed that it originates in free-free emission from an expanding isothermal shell. Initially the shell is optically thick so that the flux density is proportional to frequency squared (ν^2) at any instant and rises as time squared (t^2) as the shell expands. Eventually the shell becomes optically thin and the flux density becomes proportional to $\nu^{-0.1}$ at any given instant. The flux density then declines as the shell becomes more dilute [see Hjellming (1996) for further discussion].

It is worth considering how the centimetre and submillimetre fluxes seen in the outburst of V1494 Aql compare with the predictions of such models. At the earliest epoch (4–5 d after outburst) the submillimetre observations give a spectral index $\alpha = 2.5 \pm$

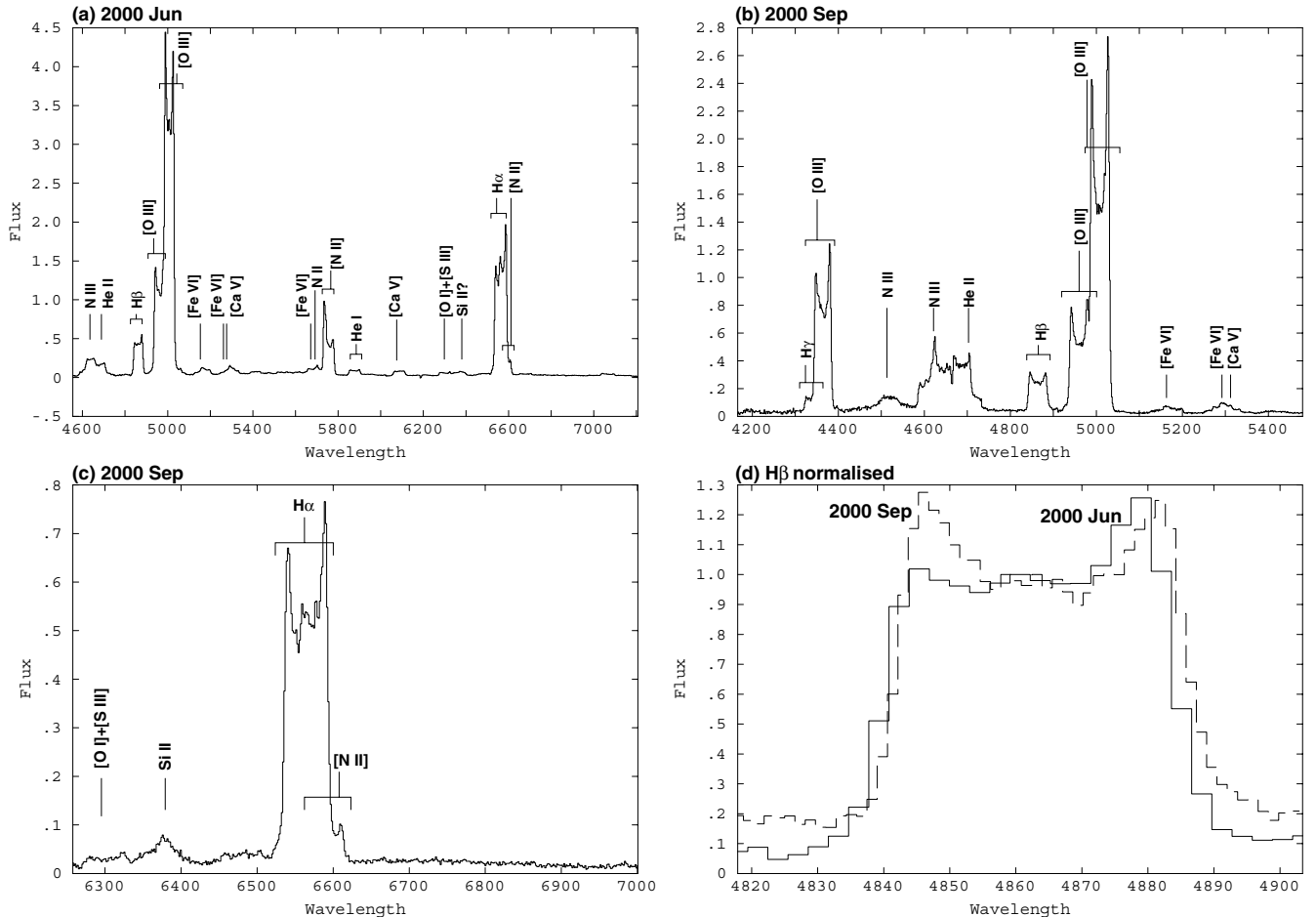


Figure 2. (a)–(c) Optical spectra with line identifications as indicated. They are not flux-calibrated. (d) $\text{H}\beta$ profiles from both dates, normalized to the central peak. The slight shift between the peaks is instrumental.

Table 2. Separation of the outermost peaks, FWHM and FWZI for the $H\beta$ profiles in Fig. 2 by wavelength and velocity equivalent. Determined from a five-component Gaussian fit in both cases.

Date	Peak separation		Full width to half-maximum		Full width to zero intensity	
	$\text{\AA}$	km s^{-1}	$\text{\AA}$	km s^{-1}	$\text{\AA}$	km s^{-1}
2000 June	35.7 ± 2.4	2200 ± 150	45.3 ± 6.0	2790 ± 370	61 ± 4	3760 ± 250
2000 September	37.7 ± 1.0	2320 ± 60	44.8 ± 3.1	2760 ± 190	66 ± 4	4070 ± 250

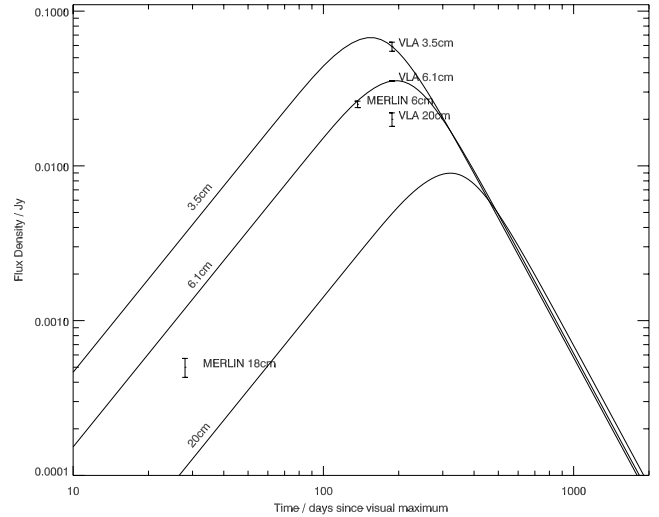
$0.5 (F_v \propto \nu^\alpha)$, which is consistent with an optically thick shell. With $\alpha = 2$ we estimate that the flux density at 6.1 cm at the same time would be $\sim 4.4 \times 10^{-3}$ mJy. We do not have an observation at 6.1 cm at this time. However, we can use the t^2 prediction of the simple models to extrapolate the estimate at 5 d and predict a 6.1-cm flux density of ~ 6.2 mJy by day 187. This falls well below the observed value of 35.4 mJy. The spectral index at this time between 6.1 and 3.5 cm is ~ 1 , suggesting that the shell was beginning to become optically thin. If that were the case, the extrapolated flux density of 6.2 mJy would represent an upper limit. Similarly extrapolating from day 5 to day 27, assuming that the emission remains optically thick, leads to a prediction for the flux density at 18 cm of around 0.02 mJy, a factor of 25 lower than that detected by MERLIN. The conclusion is that a single simple isothermal model cannot be used to fit the early submillimetre observations and the later centimetre observations simultaneously. This is in agreement with the findings of Ivison et al. (1993) for the case of V1974 Cyg.

Finally, the flux density at 6 cm on day 137, from MERLIN observations, compared with the VLA flux density at 6.1 cm on day 188, suggests a rise at $\sim t^{1.1}$. This is consistent with the simple model at turnover between the optically thin and optically thick regimes. However, by this stage MERLIN may well be resolving out some of the structure at 6 cm and hence this flux density must be considered as a lower limit. We note that changes to the $H\beta$ profiles between 2000 June and September (days 195 and 282; see Fig. 2) are consistent with the gas becoming optically thinner at optical wavelengths over that period.

We have attempted to fit the data with the two standard models for centimetre-wave radio emission from novae, the Hubble-flow model (Hjellming 1996) and the variable-wind model (Kwok 1983). As our analysis above indicates, it is impossible to obtain a satisfactory fit to the whole submillimetre–centimetre data set. Of the two models, the Hubble-flow option provides the results most consistent with the data, considering that the VLA 20-cm flux is an upper limit. The results of our best fit to the VLA data and the MERLIN data are given in Fig. 3. From this we find ejection velocities of ~ 4600 km s^{-1} and a total shell mass of $2.8 \times 10^{-4} M_\odot$. While the velocity is high even for a nova as fast as this one, the mass is typical of those determined from this kind of observation (Hjellming 1996). See Heywood et al. (2005) for a discussion of these models and our application of them to classical novae.

4.2 Oxygen line ratios

In 2000 September we observed the three [O III] lines at $\lambda\lambda 4363.5$, 4958.9 and 5006.8 $\text{\AA}$. Following Iijima & Esenoglu (2003) and using their equation (1) for the intensity ratio between the sum of the intensities of the two lines around 5000 $\text{\AA}$ and that of the line at 4363 $\text{\AA}$, we can determine the electron density n_e , using a nebula temperature of 10 700 K (Iijima & Esenoglu 2003). From Fig. 2(b) we find a ratio of 3.33 ± 0.77 . This gives a global average $n_e = (1.15 \pm 0.37) \times 10^{12} \text{ m}^{-3}$. Taking a rough ratio of the summed profiles

**Figure 3.** Radio light-curves from the Hubble-flow model, fitted to the VLA data and MERLIN data as marked. The VLA 20-cm data point should be considered an upper limit (see Table 1).

around 5000 $\text{\AA}$ to the profile at 4363 $\text{\AA}$ shows that this ratio is fairly constant across the range of velocities, suggesting little variation in density across the velocity structure (and hence presumably across the spatial structure). This value is comparable to but significantly smaller than the value of $5.2 \times 10^{12} \text{ m}^{-3}$ from Iijima & Esenoglu (2003) around the same date.

4.3 Remnant structure

In Fig. 1 we can see that the remnant consists of a partial shell, with two areas of extended emission approximately to the east and west of the central binary system. The structure is consistent with that apparent in observations of previous classical novae (Eyres et al. 1996, 2000; Heywood et al. 2005). Note that the registration between the optical and radio positions is not accurate enough to allow us to place the star precisely with respect to the radio structure. We assume that the star lies between the two radio peaks, i.e. at the centre of the partial shell.

The spectra in Fig. 2 are predominantly double-peaked, consistent with emission from a shell of material. Gill & O'Brien (1999) have demonstrated that multiply peaked spectra are consistent with substructure in the shell. However, in the case of V1494 Aql the lines are so broad that adjacent features merge, so that it is impossible to distinguish substructure in one line from the overall structure in an adjacent line. Indeed, where lines are most likely isolated (e.g. in the case of $H\beta$), the profile has little consistent substructure, and is essentially symmetric between the two wings. Table 2 gives the calculated corresponding expansion velocities for a simple shell model.

Anupama, Sahu & Mayya (2001) present spectra taken between days 3 and 22. They show dramatic changes in the $H\alpha$ profile, from a P Cygni profile on day 3, to a broad, structureless form on day 5 [also seen by Kiss & Thomson (2000)], followed by the development of structure similar to that seen in Fig. 2 by day 22. It is interesting that our JCMT observations coincide with the featureless stage of the $H\alpha$ profile, suggesting that the ejecta were in transition from optically thick to optically thin. It is also of note that the overall structure of the $H\alpha$ line profile remains similar from day 15 right through to day 195, before the blueshifted component brightens by day 285. Hjellming (1996) attributed this switch from a flat, broad spectrum to the form seen in Fig. 2(d) to a transition from continuous mass outflow to a detached shell. This would indicate a chronology in which continuous mass outflow ceases prior to day 15, followed by the expansion of the optically thick detached shell. However, we are not convinced that the changes in the profile require the end of mass outflow. For example, a equatorial ring of enhanced density would give a double-peaked profile even while mass outflow continued. We note that Iijima & Esenoglu (2003) claim evidence for continued mass ejection more than 150 d after outburst. In either case, the ejecta becomes optically thin in the visible wavelength range between days 195 and 285. This places some time constraints on future modelling of nova ejecta.

A simple Gaussian fit to the radio emission structure indicates that on day 137 the major axis is 265 ± 16 mas at position angle (east of north) 110.7 ± 4.2 , minor axis 155 ± 9 mas. This agrees well with the extent of the emission on the image (Fig. 1). In addition, the separation of the peaks of the two radio components is 120 ± 5 mas. The position angle differs from that of $\sim 140^\circ$ found around optical maximum by Kawabata et al. (2001) using spectropolarimetry. However, given the breadth of the radio features in Fig. 1, the image seems consistent with extension in the general south-east to north-west direction, as seen at early times by Kawabata et al. (2001).

Kiss & Thomson (2000) use a number of maximum magnitude–rate-of-decline relationships to find a distance of 3.6 ± 0.3 kpc. Iijima & Esenoglu (2003) find reddening amounting to $A_V = 1.8$, and consequently a modified distance of 1.6 ± 0.2 kpc. Using this latter distance estimate we can use the observed extent of the radio emission to estimate an ejection velocity. This leads to values ranging from 980 ± 80 to 2500 ± 200 km s $^{-1}$ depending on whether the major or minor axes or the peak-to-peak separation are used in the calculation. This velocity range is entirely consistent with that derived from the spectra and presented in Table 2, and confirms that the radio continuum arises from the same material as the optical lines. Thus it may be that, at this stage of the development of the radio emission (whilst the remnant is still largely optically thick at these wavelengths), it is valid to compare the extent of the remnant on the sky (Fig. 1) with velocities derived from hydrogen spectra to derive a distance to classical novae.

It has been suggested (Kiss & Thomson 2000) that the spectral line profiles are consistent with a nearly edge-on equatorial ring of the sort modelled by Gill & O’Brien (1999). We feel, given the width of the individual components, that this conclusion cannot be directly supported; we do not rule out such a structure but the spectra do not require it. None the less, we can examine the possibility that our MERLIN image (Fig. 1) represents emission from a tilted ring. Assuming that the major axis of the brightness distribution on the sky represents the diameter of an equatorial ring, we take foreshortening due to the tilt to lead to the ellipticity. Consequently the minor axis can be compared with the major axis to give an angle of inclination to the line of sight of $\sim 36.0 \pm 4.0$, which cannot be

considered edge-on. However, it does agree well with the findings of Iijima & Esenoglu (2003), who reconcile high radial velocities of absorption lines in the early decline with lower emission linewidths in the later nebular stage. They suggest that the ejecta take the form of a ring with an angle to the line of sight of 30° , also in agreement with our angle above. In this case the intrinsic velocities are twice as large as those observed and may at least partly explain why the high velocities derived from the spherically symmetric model for the radio emission presented in Section 4.1 are inconsistent with the velocities derived from the optical spectroscopy.

We note that Iijima & Esenoglu (2003) also suggest the presence of high-velocity jets in V1494 Aql. These could also be a source of the elongation in the radio emission, and the velocities of the jets are consistent with emission of the angular extent seen. However, it is difficult to see how the radio emission could be reconciled with both fast jets and an inclined ring of material; either one or the other must have caused the observed structure. Given that our MERLIN observations are contemporaneous with the nebular stage as defined by Iijima & Esenoglu (2003), it seems simplest to associate the observed structure in the radio with an inclined ring; the size-scale also agrees with this interpretation. This does not mean that jets were not present also, only that they were not visible to MERLIN.

4.4 Remnant mass

We can use some very simple assumptions to gain a model-independent estimate of the mass of the ejecta. Assuming that ejecta formed of pure, ionized hydrogen, the mass might be estimated as

$$m(H) = n_e m_H V \text{ kg} \quad (1)$$

for an electron density n_e , hydrogen atom mass m_H and ejecta volume V . We estimate V by assuming that the major and minor axes of the emission on the sky are the major and minor axes of the ellipsoidal ejecta, and that the intermediate axis is the mean of these two. Taking into account continuous mass ejection, we find the volume of the shell occupied by the ejecta on day 136 to be $V \sim 6.3 \times 10^{40}$ m 3 at a distance of 1.6 kpc. This gives $m(H) \sim 6 \times 10^{-5} M_\odot$. This is comparable to the total mass estimate of Iijima & Esenoglu (2003), but significantly less than our estimate from model fitting. If mass ejection had ended on day 15, the ejecta would only occupy around 30 per cent of this volume, with a corresponding $m(H) \sim 1.8 \times 10^{-5} M_\odot$.

5 CONCLUSIONS

We present a number of observations of the remnant of the classical nova V1494 Aql. These begin 4 d after visual maximum with submillimetre data from the JCMT, followed by centimetre-band data after 136 d (MERLIN) and 187 d (VLA) and finally optical spectroscopy on ~ 194 and ~ 284 d (UAGS). Together, these observations probe the ejecta over a period of 20 weeks, from the early ‘fireball’ stage to the final optically thin phase as the radio emission begins to fade away. Comparison between the radio image and spectroscopy is consistent with a distance of 1.6 kpc.

Both the optical spectra and the radio imaging demonstrate the early clumpy structure of the ejecta around classical novae. Most importantly, they conclusively show that simple spherical models of the ejecta development are insufficient to predict the behaviour over the time and wavelength domain sampled here (183 d, 450 μ m to 20 cm). Clearly, new models for radio emission from novae must be developed which address departures from spherical symmetry and

the clumpy nature of the ejecta. Furthermore, it is essential to investigate in detail how the predicted brightness distributions from such models would be seen by radio interferometers such as MERLIN if we are confidently to reconcile features in the radio images with structures in the ejecta. We find that the emission is consistent with either an inclined ring or jets of ejecta [as suggested by Iijima & Esenoglu (2003)] but not both; we are inclined to associate it with the former rather than the latter. We also find data consistent with mass ejection continuing for six months after the outburst right up to the ejecta starting to become optically thin in the visible regime after day 195. This continuous mass ejection needs to be tested in other novae, in particular looking for correlations with speed class.

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