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Age dissection of the vertical breathing motions in *Gaia* DR2: evidence for spiral driving

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ABSTRACT

Gaia DR2 has revealed breathing motions in the Milky Way, with stars on both sides of the Galactic mid-plane moving coherently towards or away from it. The generating mechanism of these breathing motions is thought to be spiral density waves. Here, we test this hypothesis. Using a *self-consistent*, high-resolution simulation with star formation, and which hosts prominent spirals, we first study the signatures of breathing motions excited by spirals. In the model, the breathing motions induced by the spiral structure have an increasing amplitude with distance from the mid-plane, pointing to an internal cause for them. We then show that, at fixed height, the breathing motion amplitude decreases with age. Next, we investigate the signature of the breathing motions in the *Gaia* DR2 data set. We demonstrate that, at the location with a consistently large breathing motion, the corresponding amplitude increases monotonically with distance from the mid-plane, in agreement with the model. Furthermore, we show that at the same location, the breathing motion amplitude decreases with age, again similar to what we find in the model. This strengthens the case that the observed breathing motions are driven by spiral density waves.

Key words: Galaxy: disc – Galaxy: kinematics and dynamics – Galaxy: structure – galaxies: interaction – galaxies: kinematics and dynamics – galaxies: spiral.

1 INTRODUCTION

Spirals are common features of disc galaxies in the local Universe (e.g. Elmegreen et al. 2011; Yu et al. 2018; Savchenko et al. 2020), as well as in high-redshift ($z \sim 1.8$) disc galaxies (e.g. Elmegreen & Elmegreen 2014; Willett et al. 2017; Hodge et al. 2019). The nature of spirals has been debated extensively, and, while there is broad agreement that spirals are density waves, a full consensus is yet to emerge (for detailed reviews, see e.g. Dobbs & Baba 2014; Sellwood & Carlberg 2019). A wide variety of physical mechanisms have been proposed for exciting spirals, including bars (e.g. Athanassoula, Romero-Gómez & Masdemont 2009; Athanassoula et al. 2010; Salo et al. 2010; Athanassoula 2012; Efthymiopoulos, Harsoula & Contopoulos 2020), tidal encounters (e.g. Toomre & Toomre 1972; Dobbs et al. 2010), swing amplification of noise (Goldreich & Lynden-Bell 1965; Julian & Toomre 1966; Toomre 1981), giant molecular clouds (D’Onghia, Vogelsberger & Hernquist 2013), other spirals (Masset & Tagger 1997), and recurrent groove modes (Sellwood & Lin 1989; Sellwood & Kahn 1991; Sellwood 2012; Sellwood & Carlberg 2019). Star formation plays a pivotal role in the persistence of spirals, by cooling the stellar disc and facilitating the generation of fresh spirals (Sellwood & Carlberg 1984), while the interstellar gas helps spiral density waves to survive for longer (Ghosh & Jog 2015, 2016). The vital dynamical role of spiral arms in transporting angular momentum

(Lynden-Bell & Kalnajs 1972), as well as in the radial mixing of stars without heating (Sellwood & Binney 2002; Roškar et al. 2008; Schönrich & Binney 2009) has kept the study of the formation and evolution of spirals of continuing interest.

The Milky Way is a barred galaxy (Weinberg 1992) that also hosts spiral structure (e.g. Gerhard 2002, also see Vallée 2005, 2008). The presence of large-scale, non-zero mean vertical motions of Solar Neighbourhood stars has been reported in various surveys, for example, using main-sequence stars in the Sloan Extension for Galactic Understanding and Exploration (SEGUE) survey (Widrow et al. 2012), F-type stars in the Large Sky Area Multi-Object Fibre Spectroscopic Telescope (LAMOST) survey (Carlin et al. 2013), and red clump stars in the Radial Velocity Experiment (RAVE) data (Williams et al. 2013). These vertical motions display two distinct types: one in which stars on either side of the mid-plane move coherently in the same direction along the perpendicular to the disc (bending motion) and the other in which stars move coherently in the opposite direction, thereby compressing towards the mid-plane or expanding away from it (breathing motion). In addition, Widrow et al. (2012) reported evidence for wave-like north–south asymmetries in the number counts of stars in the Solar Neighbourhood, which were further explored by Yanny & Gardner (2013).

The *Gaia* mission (Gaia Collaboration 2016, 2018a) has revolutionized the study of Galactic archaeology by measuring the stellar kinematics of stars in the Solar Neighbourhood with an unprecedented precision. The stellar kinematic study from the second *Gaia* Data Release (hereafter *Gaia* DR2) revealed rich kinematic

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substructure in the phase-space distribution of Solar Neighbourhood stars (Antoja et al. 2018), as well as the presence of large-scale non-zero vertical motions ($\sim 10 \text{ km s}^{-1}$ in magnitude), and associated bending and breathing motions (Gaia Collaboration 2018b). Using the RAVE and the Tycho–Gaia astrometric solution (TGAS) catalogue, Carrillo et al. (2018) showed the presence of both bending and breathing motions in the Solar Neighbourhood. Bennett & Bovy (2019) revisited the north–south asymmetry in the Solar Neighbourhood stars using the position–velocity measurements from *Gaia* DR2, and confirmed the presence of periodic over- and underdensities of stars in the vertical distribution, regardless of their colours. For an axisymmetric potential, the bulk radial and vertical motions should be zero (e.g. Binney & Tremaine 2008). Therefore, the question arises what mechanism(s) can induce these non-zero bulk vertical motions in our Galaxy.

Theoretical efforts to understand the cause of vertical motions have explored both external and internal mechanisms. An interaction with a satellite or a dark matter subhalo can excite large-scale coherent vertical bending motions (e.g. see Hunter & Toomre 1969; Araki 1985; Mathur 1990; Weinberg 1991; Gómez et al. 2013; Widrow et al. 2014; D’Onghia et al. 2016; Chequers, Widrow & Darling 2018). However, bending motions can be produced by purely internal mechanisms, such as the bar (Khoperskov et al. 2019) or the warp (Khachatryan et al. submitted). On the other hand, internal causes such as spiral density waves (Debattista 2014; Faure, Siebert & Famaey 2014; Monari, Famaey & Siebert 2016) or the Galactic bar (Monari, Famaey & Siebert 2015) as well as external mechanisms, such as a fly-by encounter with a satellite or passing of a dark matter sub-halo (Widrow et al. 2014), have been proposed for exciting the breathing motions. Earlier studies by Siebert et al. (2011, 2012) showed that spirals and bars can drive a large-scale gradient in the bulk radial velocity in the Milky Way. Strong spirals drive large-scale vertical breathing motions ($|\langle v_z \rangle| \sim 5\text{--}20 \text{ km s}^{-1}$) as first shown by Faure et al. (2014). Their semi-analytic models assumed a thin spiral structure with a small vertical extent. Their analytic treatment found that the breathing motions vary with distance from the mid-plane as $\tanh(z/z_0)\text{sech}^2(z/z_0)$, where z_0 is the scale height of the spiral. This peaks at $\sim 0.7z_0$. At larger heights, this decreases as $\exp(-2z/z_0)$ (see equations 10 and 12 in Faure et al. 2014). The self-consistent simulations of Debattista (2014), instead, produced vertically extended spirals, which resulted in vertical motions that increase with distance from the mid-plane. The relative sense of the bulk motions, whether compressing or expanding changes across the corotation resonance (hereafter CR) of the spiral. Inside the CR, the vertical motions are compressive behind the peak of the spiral and expanding ahead of the spiral’s peak. Outside the CR, the compressive and expanding breathing motions reverse their sense of occurrence (Debattista 2014; Faure et al. 2014).

In this paper, we continue to explore the possibility that the breathing motions observed in *Gaia* DR2 are caused by the Milky Way’s spiral structure. Using a high-resolution, self-consistent, star-forming simulation of a Milky Way-like galaxy, we study the characteristics of the breathing motions driven by spiral structure. Previous studies of breathing motions either used perturbation theory (Monari et al. 2016) and test particle simulations (Faure et al. 2014) or used self-consistent simulations but lacked the high-mass resolution necessary to sub-sample the stellar populations (Debattista 2014). The novelty of our paper, aside from comparing directly with the *Gaia* DR2 data, lies in the fact that our high-resolution model allows us to explore how different populations participate in the breathing motions.

The paper is organized as follows: Section 2 provides the details of the simulation, including quantifying the spiral structure present

at our chosen snapshot, which, for simplicity, we refer to as ‘the model’. Section 3 presents the breathing motions of the model. Section 4 provides the details of a novel technique to detect/identify breathing motions from the variation of the bulk vertical motions. Section 5 compares the vertical kinematic signatures of Milky Way stars with the trends obtained in the model. Section 6 discusses the *unique* fingerprint of spiral-driven vertical breathing motions, while Section 7 summarizes our main conclusions.

2 STAR-FORMING SIMULATION

We use an N -body + smooth particle hydrodynamics (SPH) simulation with gas and star formation to motivate some of the analysis of the *Gaia* DR2 data. This is a higher mass resolution version of the models described in our previous papers (e.g. Roškar et al. 2008; Loebman et al. 2016). This model starts with a gas corona in pressure equilibrium with a co-spatial Navarro–Frenk–White (Navarro, Frenk & White 1996) dark matter halo which constitutes 90 per cent of the mass. The dark matter halo has a virial radius $r_{200} \simeq 200 \text{ kpc}$ and a virial mass $M_{200} = 10^{12} M_\odot$. Gas velocities are given a spin of $\lambda = 0.065$ (Bullock et al. 2001), with specific angular momentum proportional to radius. Both the gas corona and the dark matter halo consist of 5×10^6 particles; gas particles have softening $\epsilon = 50 \text{ pc}$, while that of dark matter particles is $\epsilon = 100 \text{ pc}$. All stars form out of cooling gas, inheriting the softening of the parent gas particle. We evolve the simulation for 13 Gyr with GASOLINE (Wadsley, Stadel & Quinn 2004; Wadsley, Keller & Quinn 2017). As gas cools it settles into a disc; once the gas density exceeds 0.1 cm^{-3} and the temperature drops below $15\,000 \text{ K}$, star formation and supernova blast wave feedback commence (Stinson et al. 2006). Supernovae feedback couples 40 per cent of the 10^{51} erg per supernova to the interstellar medium as thermal energy. Gas mixing uses turbulent diffusion as described by Shen et al. (2010).

We use a base time-step of $\Delta t = 5 \text{ Myr}$ with time-steps refined such that $\delta t = \Delta t/2^n < \eta\sqrt{\epsilon/a_g}$, where we set the refinement parameter $\eta = 0.175$. We set the opening angle of the tree-code gravity calculation to $\theta = 0.7$. Gas particle time-steps also satisfy the condition $\delta t_{\text{gas}} = \eta_{\text{courant}} h / [(1 + \alpha)c + \beta\mu_{\text{max}}]$, where $\eta_{\text{courant}} = 0.4$, h is the SPH smoothing length set over the nearest 32 particles, α and β are the linear and quadratic viscosity coefficients, and μ_{max} is described in Wadsley et al. (2004). Star particles represent single stellar populations with a Miller–Scalo initial mass function.

2.1 Spiral structure in the model

We use a snapshot from this model at $t = 12 \text{ Gyr}$. We choose this snapshot strictly because it allows a good division of its stellar populations, not because of any special properties of the model at this time. While the simulation we use has not been published before, the general evolution of the spiral structure in this class of models is explored in Roškar et al. (2012).

In order to facilitate comparison with the Milky Way, we quantify the spiral amplitude in the model to show that the spirals in the model are not unreasonably strong. We also quantify the strength of the spiral as a function of stellar age. We divide the stars into three broad age bins: a young stellar population (with ages between 0 and 4 Gyr), an intermediate-age stellar population (with ages between 4 and 8 Gyr), and an old stellar population (with ages between 8 and 12 Gyr). The top panels of Fig. 1 show the surface density of stars in the face-on, (x, y) , view for the three age populations. The bottom

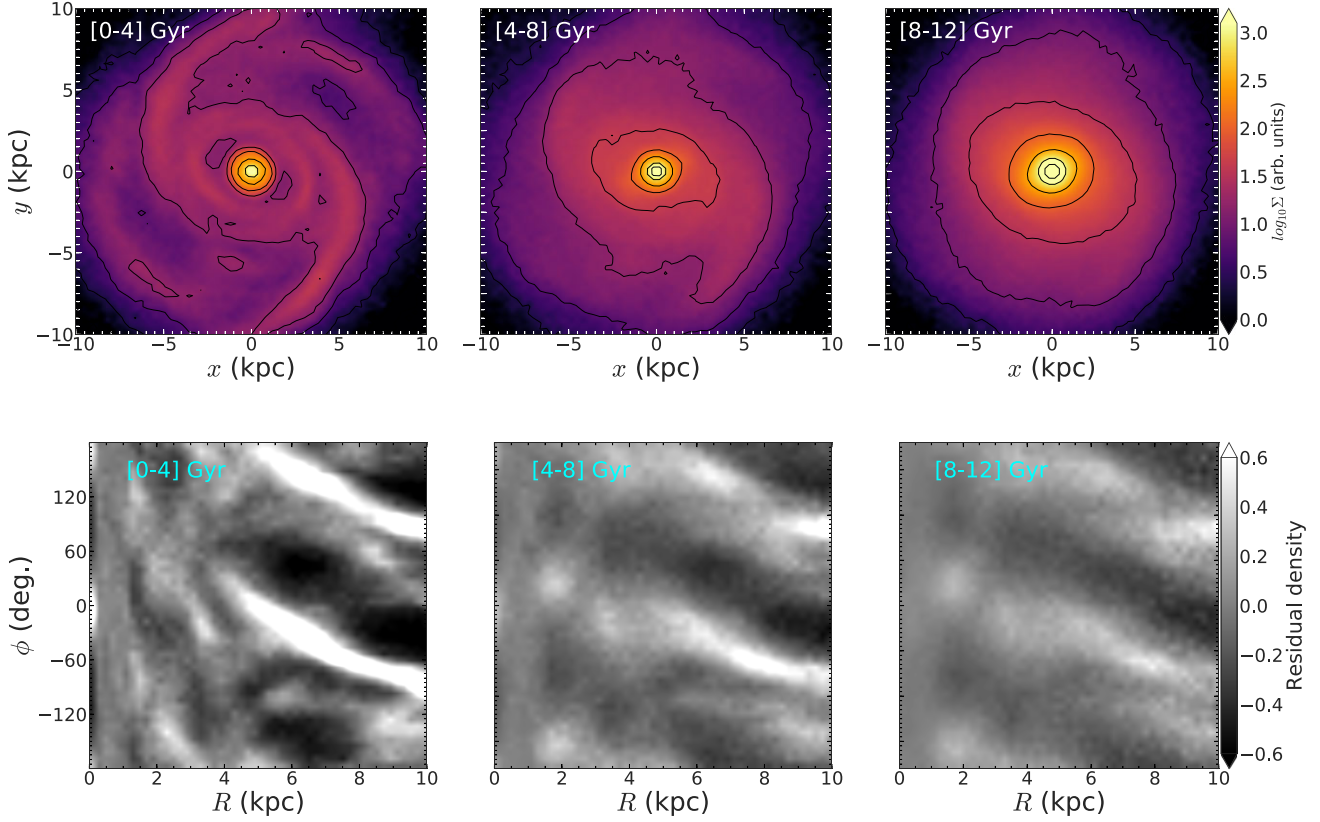


Figure 1. Star-forming model: distribution of stellar surface density (top panels), and the residual surface density (bottom panels), calculated using equation (1) at $t = 12$ Gyr for (left to right) young (ages = 0–4 Gyr), intermediate (ages = 4–8 Gyr), and old (ages = 8–12 Gyr) populations. Contours of constant surface density (black solid lines) are overlaid on the surface density maps. Unambiguous spiral structure is present in the young and intermediate age populations, and the residuals reveal that the old population also supports the spiral. The sense of rotation is towards increasing ϕ .

panels show the residual surface density, $\tilde{\Sigma}(R, \phi)$, calculated as

$$\tilde{\Sigma}(R, \phi) = \frac{\Sigma(R, \phi) - \Sigma_{\text{avg}}(R)}{\Sigma_{\text{avg}}(R)}, \quad (1)$$

where $\Sigma_{\text{avg}}(R)$ is the average density at radius R , and the averaging is carried out over the whole azimuthal range.

The density plots of Fig. 1 show that the youngest population exhibits a very prominent two-armed spiral structure, which is also present in the intermediate-age population, but is indistinct in the old population. However, the residual density plots show that the $m = 2$ spiral is present even in the old population, albeit weaker. This therefore is a genuine density wave which propagates in all stellar populations. The estimated pitch angle, γ , of the spiral in this snapshot is $\sim 33^\circ$. In comparison, the Milky Way's spirals are tighter, having a pitch angle $\gamma \sim 10^\circ$ (Siebert et al. 2012).

To quantify the strength of the spirals, we measure the radial profiles of the Fourier moments of the surface density, defined as

$$A_m/A_0(R) = \frac{\sum_i m_i e^{im\phi_i}}{\sum_i m_i}, \quad (2)$$

where A_m is the coefficient of the m th Fourier moment of the density distribution, m_i is the mass of the i th particle, and ϕ_i is its cylindrical angle. The peak of the Fourier $m = 2$ amplitude, $(A_2/A_0)_{\text{max}} \simeq 0.19$ for our model. In comparison, Rix & Zaritsky (1995) found that, for their sample of galaxies with $m = 2$ spirals, the peak value of A_2/A_0 varies in the range 0.15–0.6, with ~ 28 per cent of their sample having $(A_2/A_0)_{\text{max}} \leq 0.2$. The spiral structure in the model therefore

is sufficiently realistic to provide a basis for comparing the kinematic signature of the spiral in the model with that in the Milky Way. The amplitudes for the $m = 2$ and $m = 4$ Fourier harmonics for all three stellar populations at $t = 12$ Gyr are shown in Fig. 2. The spiral has a significant $m = 2$ amplitude at $4 \leq R/\text{kpc} \leq 8$. While the amplitude decreases with stellar age, all stellar populations take part in the spiral structure, as opposed to *material spirals* present in only the young stellar populations (for detailed discussion, see e.g. Binney & Tremaine 2008; Dobbs & Baba 2014). The $m = 4$ amplitude is likewise strongest in the young population, and is almost zero for the old population. Thus, the spiral in the old population varies more gently in azimuth while that in the young population is more sharply delineated, as can be seen in Fig. 1.

We also calculate the residual density distribution $\delta\Sigma(\phi, z)$ within a radial annulus of width ΔR as

$$\delta\Sigma(\phi, z; \Delta R) = \frac{\Sigma(\phi, z; \Delta R) - \Sigma_{\text{avg}}(z; \Delta R)}{\Sigma_{\text{avg}}(z; \Delta R)}, \quad (3)$$

where $\Sigma_{\text{avg}}(z; \Delta R)$ is the average density, and the averaging is carried out over the whole azimuthal range within the radial annulus of radial width ΔR , at a fixed height z . Fig. 3 shows the resulting residual maps in the (ϕ, z) plane at different annuli of $\Delta R = 1$ kpc. At each annulus, the peak density in the mid-plane ($z = 0$) is ahead (relative to the sense of rotation) of the peak density at larger heights from the mid-plane. Also, the density distribution at fixed R and z is azimuthally asymmetric relative to the peak density, in agreement with the findings of Debattista (2014). The spiral extends to the

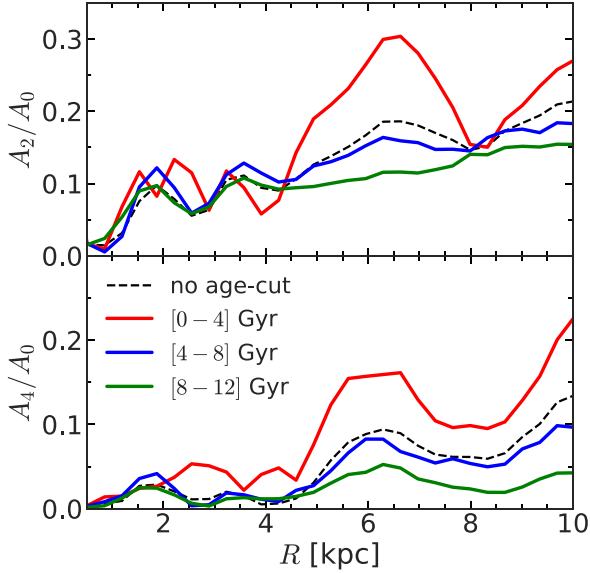


Figure 2. Star-forming model: radial profile of $m = 2$ (top panel) and $m = 4$ (bottom panel) Fourier moments of the surface density for the three stellar populations at $t = 12$ Gyr. The nearly zero A_4 in the old population indicates that the spiral in these stars is more gently varying with azimuth than in younger populations. Note the change of vertical axis scale between the two panels. The black dashed lines denote the Fourier moments for all stars.

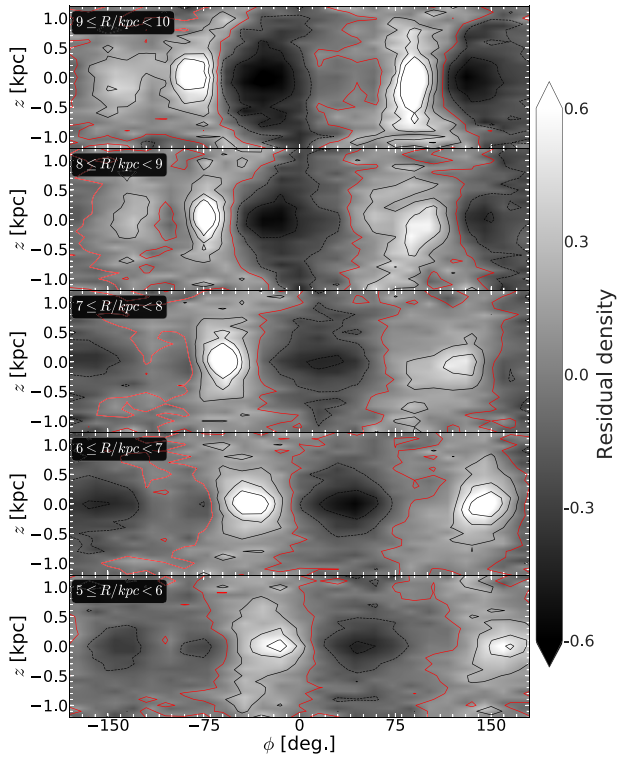


Figure 3. Star-forming model: residual surface density, $\delta\Sigma(\phi, z; \Delta R)$, calculated using equation (3), shown in the (ϕ, z) plane at five different radial annuli of width $\Delta R = 1$ kpc. The contours indicate constant residual surface density, with the red contours denoting $\delta\Sigma(\phi, z; \Delta R) = 0$. The sense of rotation is towards increasing ϕ .

full vertical extent, rather than being limited to close to the mid-plane. This behaviour was also found in pure N -body simulations by Debattista (2014), and will be important for our subsequent analysis.

3 BREATHING MOTIONS INDUCED BY SPIRALS

We start by studying the breathing motions induced by the spirals in the star-forming model in order to explore how the breathing motions vary between different age populations. The self-consistent simulation of Debattista (2014) showed that the amplitudes of the breathing motion increases with height from the mid-plane; here, we extend that earlier work but exploring how the amplitudes change as a function of both height and stellar age.

We define the breathing velocity, V_{breath} , as (Debattista 2014; Widrow et al. 2014; Gaia Collaboration 2018b):

$$V_{\text{breath}}(x, y) = \frac{1}{2} [\langle v_z(x, y, \Delta z) \rangle - \langle v_z(x, y, -\Delta z) \rangle], \quad (4)$$

where $\langle v_z(x, y, \Delta z) \rangle$ is the mean vertical velocity at position (x, y) in the galactocentric Cartesian coordinate system, averaged over a horizontal layer of thickness Δz (for details, see Gaia Collaboration 2018b). We choose $\Delta z = 400$ pc, and calculate V_{breath} in three such vertical layers: $|z| = [0, 400]$ pc, $|z| = [400, 800]$ pc, and $|z| = [800, 1200]$ pc. In this definition of V_{breath} , the average in a particular layer is carried out before the difference between the two sides is taken. In the observational data, this helps reduce the effect of any selection function differences between the two sides of the disc. This also removes the trivial effect of bending waves, which would otherwise give rise to a spurious signature by displacing stars to one side of the mid-plane. Equation (4) implies that when $V_{\text{breath}} > 0$ then stars are coherently moving away from the mid-plane (expanding breathing motion), while $V_{\text{breath}} < 0$ when stars are moving towards the mid-plane (compressing breathing motion). For the sake of comparison, we also consider the bending velocity:

$$V_{\text{bend}}(x, y) = \frac{1}{2} [\langle v_z(x, y, \Delta z) \rangle + \langle v_z(x, y, -\Delta z) \rangle]. \quad (5)$$

The star-forming model, which is not externally perturbed, lacks substantial coherent bending waves, so the model's bending velocities are not shown here.

3.1 Bulk vertical motions of the different age populations

We examine the star-forming model's vertical motions as a function of stellar age. The distribution of the bulk vertical motions in the meridional, (R, z) , plane, calculated at four different azimuthal wedges, for the three stellar populations at $t = 12$ Gyr, are shown in Fig. 4. Large-scale, non-zero vertical motions are present in all three populations. The absolute values of the mean vertical velocity, $|\langle v_z \rangle|$, are small, ~ 5 km s $^{-1}$, but non-zero, increasing with height from the mid-plane. There are not enough stellar particles in the young population at large heights to measure a reliable mean vertical velocity for them there. The intermediate-age population exhibits a large $|\langle v_z \rangle|$ at larger heights; $|\langle v_z \rangle|$ gets somewhat weaker for the old stellar population. All three populations are coherently moving away from, or towards, the disc mid-plane indicating breathing motions predominate in all stellar populations. There are also bending motions present, but these are weaker compared to the breathing motions. These bending motions are studied in detail elsewhere (Khachatryan et al. submitted). All three age populations exhibit prominent bulk vertical breathing motions in the (R, z) plane.

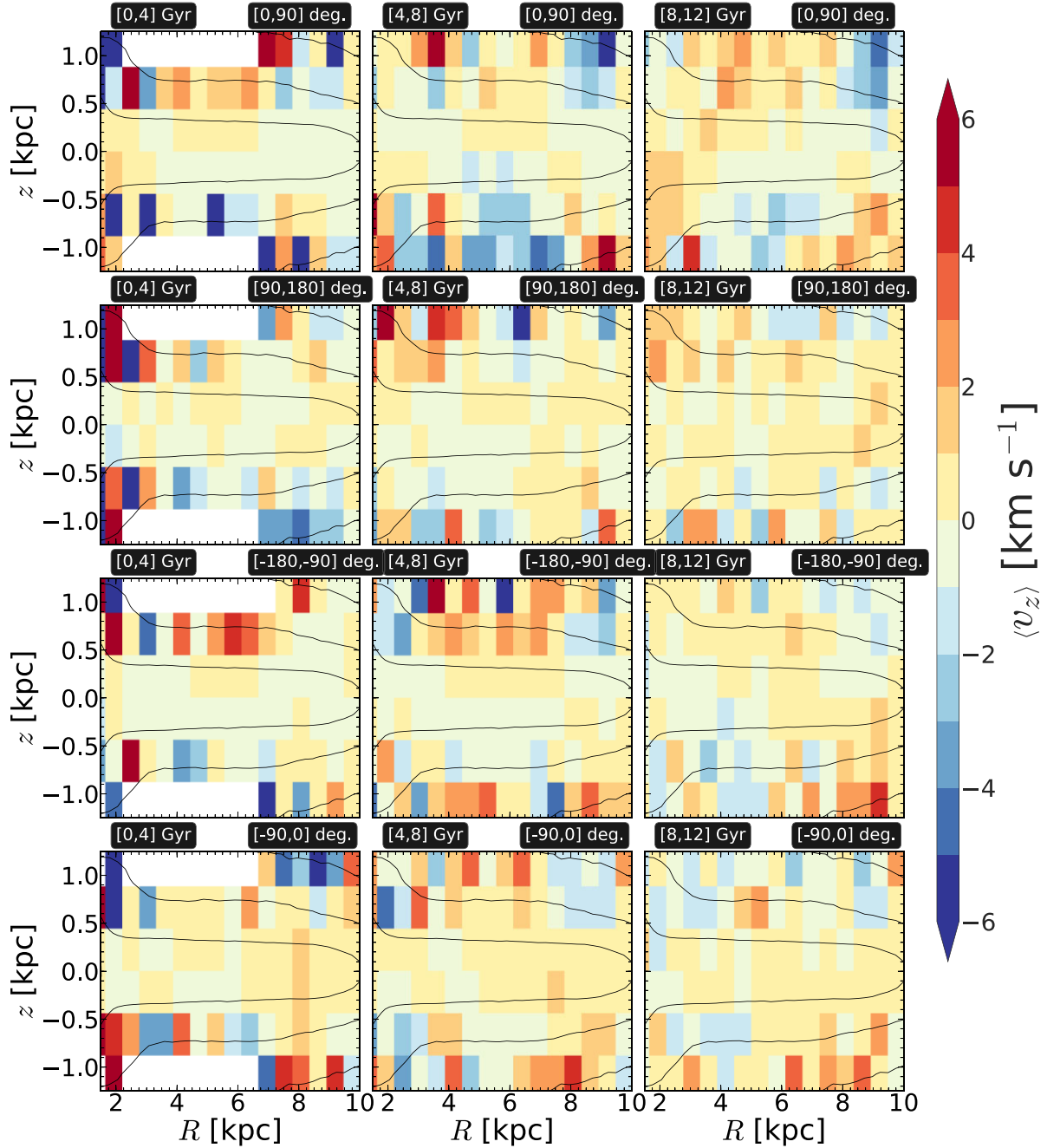


Figure 4. Star-forming model: distribution of bulk vertical velocity, $\langle v_z \rangle$, in the meridional, (R, z) , plane, for the three different stellar populations at $t = 12$ Gyr. Left to right: stars age = 0–4 Gyr (young), = 4–8 Gyr (intermediate), and 8–12 Gyr (old). Top to bottom: stars at different azimuthal wedges are shown. The age bins and the azimuthal wedges are listed above each sub-plot. Contours of the total density are overlaid in each panel.

However, the azimuthal wedges used in Fig. 4 are large, therefore the coincidence of strong breathing motion with the spiral structure is not readily apparent in this figure. This is further demonstrated in Section 4.

We then select stars in the radial annulus $5 \leq R/\text{kpc} \leq 7$, calculate $\langle v_z \rangle$ in the (ϕ, z) plane for the different stellar populations, and plot these in Fig. 5. Non-zero bulk vertical motions are evident in all the stellar populations, although the young population particle numbers are not large enough at the uppermost slice to measure a meaningful $\langle v_z \rangle$. These results indicate breathing motions are present in this annulus. The vertical motions switch from compressive to

expanding at $\phi \simeq 0^\circ$, which Fig. 1 shows is the location of the peak density in the youngest population. The amplitude of the vertical motions decreases with age, with the old population having velocities roughly half those of the intermediate population.

3.2 Breathing motions of the different age populations

We therefore compute the breathing velocities for stars of different ages in three vertical slices: $|z| = [0, 400]\text{ pc}$, $|z| = [400, 800]\text{ pc}$, and $|z| = [800, 1200]\text{ pc}$. We calculate V_{breath} using equation (4) separately for each age population. Fig. 6 shows V_{breath} in the $(x,$

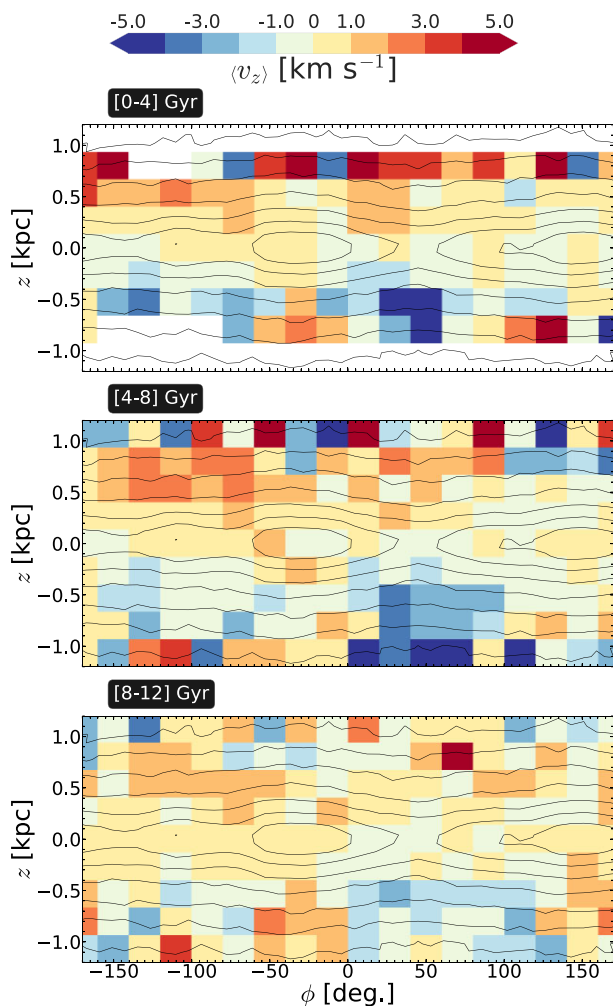


Figure 5. Star-forming model: distribution of $\langle v_z \rangle$ in the (ϕ, z) plane, for stars in the radial annulus $5 \leq R/\text{kpc} \leq 7$ at $t = 12$ Gyr. Top: the young stellar population, middle: the intermediate-age stellar population, and bottom: the old stellar population. The presence of large-scale, non-zero bulk vertical velocities for stars in all three populations is visible, and are coherent across age. The vertical motions are smallest in the old population. The sense of rotation is towards increasing ϕ .

y) plane whereas Fig. 7 shows V_{breath} in the (R, ϕ) plane. Strong breathing motions are present in all three age populations; V_{breath} has small amplitude near the mid-plane and increases with distance from it. This trend is present for all ages; at the largest heights the intermediate-age population has the strongest V_{breath} amplitude, but the youngest population does not reach these heights so can only be compared in the intermediate layer. The intermediate population at the largest height shows that the compressive motions are closely associated with the peak density of the spiral in the outer region (excluding the central weak bar).

In summary, we find large-scale, non-zero vertical motions present both in the meridional plane as well as in the (ϕ, z) plane at a fixed radial annulus. These non-zero bulk vertical motions appear preferentially at azimuthal locations that contain the strongest spiral structure. The resulting breathing velocity, V_{breath} , is small near the disc mid-plane, and increases with height, as found also by Debattista (2014). The breathing velocity in our model is largest for the intermediate age population at the largest heights, because the

young population does not reach such heights; in the middle layer, the young population has the largest V_{breath} , while the old population has the smallest.

4 A NOVEL TECHNIQUE TO DETECT THE BREATHING MOTIONS

In the previous section, we measured the distribution of breathing velocity in the (x, y) plane using equation (4). This requires calculating, at a particular spatial location, the mean vertical velocity, $\langle v_z \rangle$, of stars in both the upper and the lower layers (of width Δz) of the disc, and then taking the difference of these two mean vertical velocities. For the simulated galaxy model, we have the whole spatial coverage so that we can calculate quite robustly the spatial distribution of the resulting breathing velocities for stars at different vertical layers. However, for the Milky Way, this may not be the case due to the limitations in the spatial coverage and incompleteness of *Gaia* DR2.

Here, we develop a new technique for identifying breathing motions from the distribution of the mean vertical velocities ($\langle v_z \rangle$) as a function of height from the disc mid-plane ($z = 0$). First, we demonstrate this technique on our star-forming model, and then apply it to the *Gaia* DR2 data.

Above we showed that the breathing velocities increase away from the mid-plane. Therefore to leading order, the variation of $\langle v_z \rangle$ with height can be fitted with a straight line. Breathing motions result in best-fitting straight lines to the height variation of the mean vertical velocity that has a significantly *non-zero* slope. If the slope is positive, then it indicates an expanding (positive) breathing motion, while if the slope is negative then it indicates the presence of a compressing (negative) breathing motion.

To test this, we divide the stellar particles into two age populations: younger (1–6 Gyr) and older (6–12 Gyr) stellar particles. We exclude the very young stellar particles (age less than 1 Gyr) because of their limited vertical coverage. Fig. 8 shows a few examples of fitting a straight line to the variation of $\langle v_z \rangle$ with height. The straight line is fitted via the SCIPY package CURVEFIT that uses the Levenberg–Marquardt algorithm. The non-zero slope indicates breathing motions with amplitude increasing with height from the mid-plane. The (absolute) value of the slope is larger for the younger stellar population than for the older stellar population. This implies that the breathing motions are stronger in the younger population than the older populations, in agreement with the findings of the previous section. We note that, at certain locations, there is a clear difference in the zero-point of the different age groups. When studied in detail, we found that these zero-point offsets arise because of weak bending waves, which have different amplitudes in the different age groups (Khachaturyants et al. submitted).

We compute the error on the slope [$\epsilon(\text{slope})$] and plot this in the middle row of Fig. 9. For uniform comparison, we impose a simultaneous quality cut of $\text{slope}/\epsilon(\text{slope}) > 3$ for both age populations, and choose only those bins where this quality cut is met. The resulting distribution of the best-fitting slopes is shown in the bottom panels of Fig. 9. Clear signatures of breathing motions (with amplitude increasing with height) are seen. The best-fitting slopes for the younger population (1–6 Gyr) are, in general, larger than those for the older population (6–12 Gyr).

Lastly, we investigate how the slope varies as a function of the azimuthal angle for the young and the old populations relative to the density variation. To achieve this, we first choose a radial extent ranging from 6 to 8 kpc (i.e. outside the CR of the spiral), where the model exhibits a strong spiral. It is evident from Fig. 1 (bottom panels) that the azimuthal locations of the density peaks

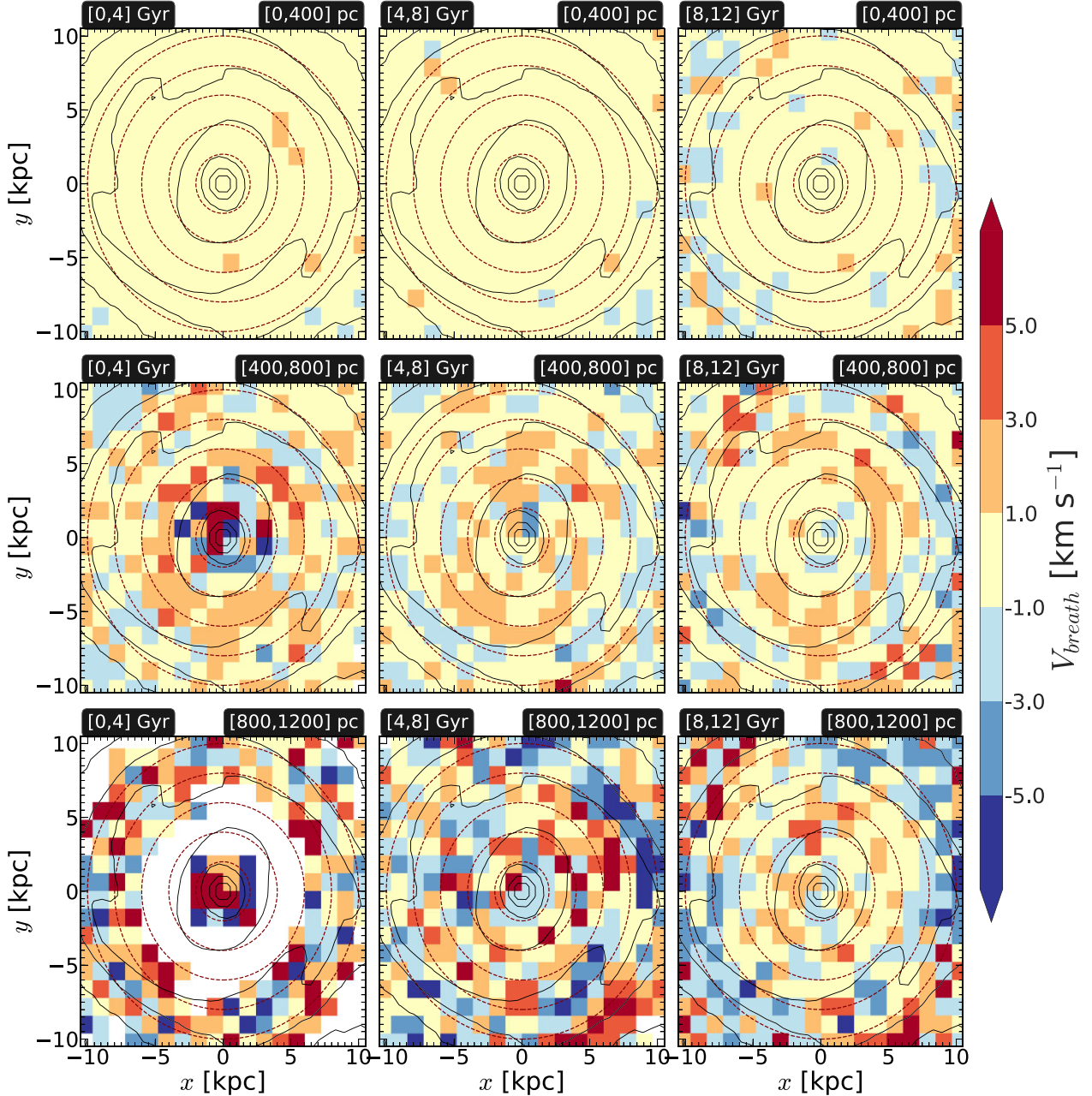


Figure 6. Star-forming model: breathing velocity, V_{breath} , for stars of different ages at different vertical distances from the mid-plane, shown at $t = 12$ Gyr. From left to right: stars with ages 0–4 Gyr (young), 4–8 Gyr (intermediate), and 8–12 Gyr (old), respectively, whereas from top to bottom: stars at $|z| = [0, 400]$ pc, at $|z| = [400, 800]$ pc, and at $|z| = [800, 1200]$ pc, respectively. Contours of *total* surface density are overlaid in each panel. The dashed circles (in maroon) indicate galactic radii ranging from 2 to 10 kpc, in 2 kpc steps.

vary as a function of radius in the chosen radial extent, because of the spiral’s pitch angle. Therefore, to obtain a stronger signal of the slope, we rotate the stars in different 1 kpc-width radial bins, in such a way that the density peaks in our chosen radial extent coincide. Then, we recalculate the slope of the breathing velocity as a function of the rotated azimuthal angle (ϕ') for the young and the old stellar populations. This is shown in Fig. 10. We also calculate the total residual surface density [$\tilde{\Sigma}(R, \phi')$] as a function of the rotated azimuthal angle in this chosen radial extent. As evident from Fig. 10, the locations of the larger slopes of the breathing motions coincide with the azimuthal locations of the density minima while the

azimuthal locations of the maximum density exhibit smaller values of the slope. In other words, the amplitudes of the breathing motions are larger in the inter-arm region. This trend is consistent with the earlier findings of Debattista (2014) that stars on the expanding side of the spiral always exhibit larger breathing motions when compared with those for the compressing side of the spiral. Debattista (2014) attributed this behaviour to the more abrupt density variation as stars leave the spirals compared with when they enter them. [The test particle and semi-analytic calculations of Faure et al. (2014) did not exhibit this behaviour because the spiral perturbation they used had a purely $m = 2$ multiplicity.] Furthermore, the slope of the young stellar

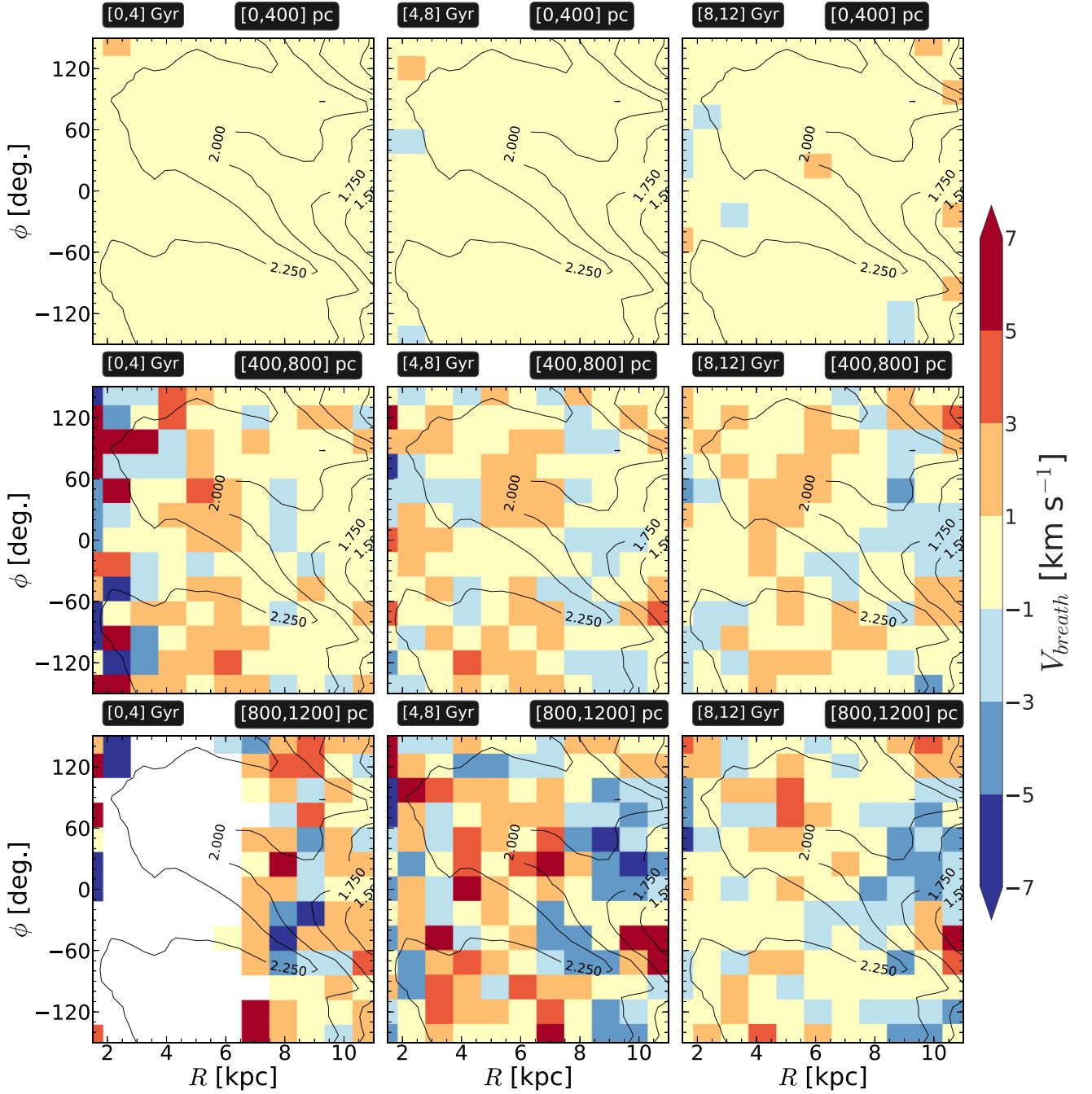


Figure 7. Star-forming model: Same as Fig. 6, but shown in the (R, ϕ) plane. The age bins and the vertical layers are presented above each panel. Contours of *total* surface density (in log scale) are overlaid in each panel. Large $|V_{\text{breath}}|$ is associated with density peaks and troughs of the spiral.

population is larger than the slope for the older stellar population, thereby demonstrating that indeed the breathing motion is stronger in the young stellar population than in the old stellar population.

5 COMPARISON WITH VERTICAL MOTIONS IN THE MILKY WAY

We now search for these signatures of spiral-driven breathing motions in the *Gaia* DR2 data as a function of stellar ages, and compare with the results obtained from the star-forming model.

5.1 Sample selection from the Sanders and Das data set

We use the publicly available sample of stellar ages of Sanders & Das (2018) (hereafter SD18), which provides a catalogue of distance, mass, and age for approximately 3 million stars in the *Gaia* DR2. This sample allows us to study the vertical motions for different age populations. *Gaia* DR2 suffers from distance biases from the parallax measurements (for detailed discussion, see e.g. Bailer-Jones et al. 2018; Lindegren et al. 2018; Luri et al. 2018; Schönrich, McMillan & Eyer 2019). To account for these biases, especially when the parallax measurements are of the same order as the associated uncertainties,

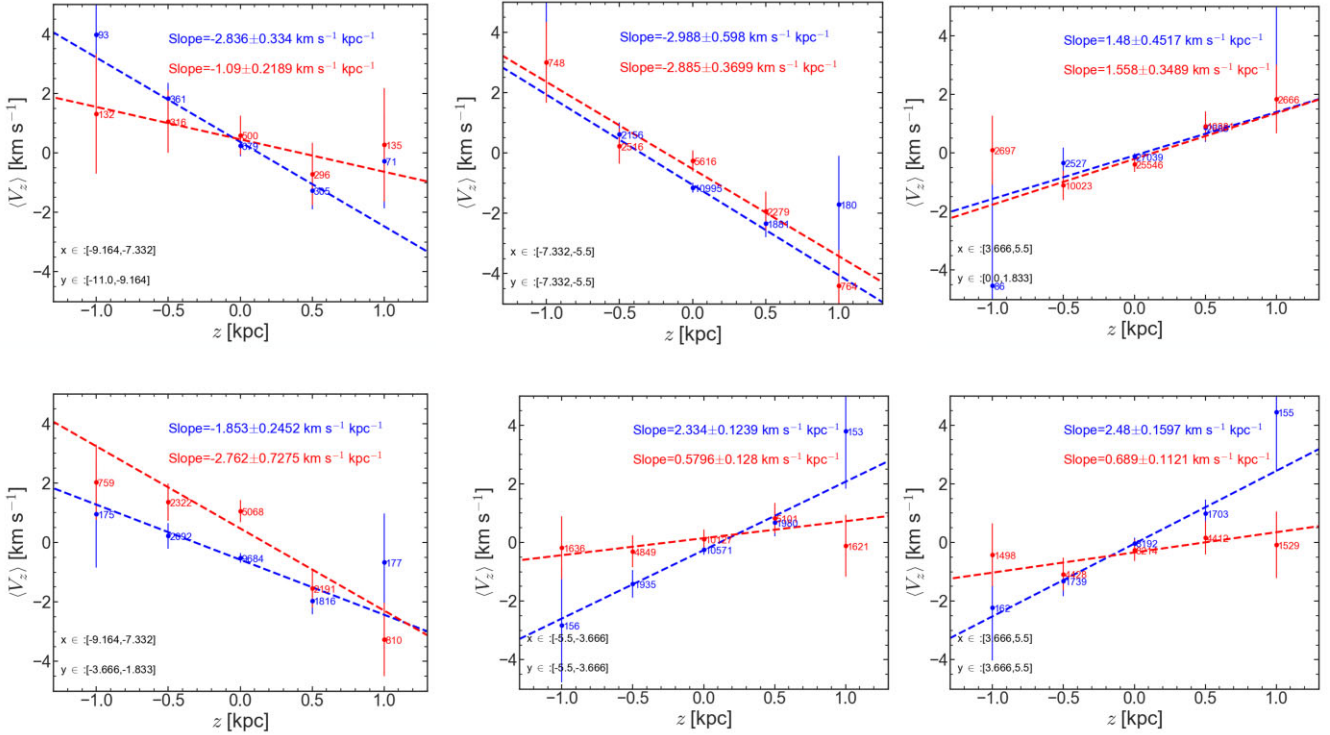


Figure 8. Star-forming model: a few examples of fitting straight lines to the variation of mean vertical velocity ($\langle v_z \rangle$) with height (z) for the younger (shown in blue) and the older population (shown in red) are shown. The resulting best-fitting slopes and the associated errors are indicated in each sub-panel. The values listed near each point denote the total number of stellar particles present in that bin. The (absolute) values of the best-fitting slope for the younger population are, in general, larger than that for the older population.

SD18 used a full Bayesian framework to calculate the distance of a star while using different photometric and spectroscopic parameters from a number of ground-based surveys (APOGEE, *Gaia*-ESO, GALAH, LAMOST, RAVE, and SEGUE), and the astrometric information from *Gaia* DR2, simultaneously as priors (for details, see section 2 of Sanders & Das 2018). Similarly, the probability distributions for mass, and ages for each star were calculated by SD18 via a Bayesian framework based on the photometric, spectroscopic, and astrometric quantities. This data set includes the six-dimensional Galactocentric cylindrical position–velocity coordinates, with the Sun placed at a Galactocentric radius $R_0 = 8.2$ kpc and a height above the plane $z_0 = 15$ pc (Bland-Hawthorn & Gerhard 2016), and the values of the Solar peculiar motions from Schönrich, Binney & Dehnen (2010).

Each stellar entry in the SD18 data set is characterized by three flags that, depending on the nature of the analysis, can aid in filtering out problematic data. The FLAG entry characterizes all issues related to the quality of the input data (spectroscopy, astrometry, photometry) from the above listed surveys and of the output data from the SD18 Bayesian frameworks (mass, age). If there are no issues in the input or output data, then FLAG = 0. The DUPLICATED flag specifies if stellar entries are duplicates due to a crossmatch within a single survey or between multiple surveys. Within a single survey, the duplicate with the smallest vertical velocity uncertainty is kept, while duplicates between surveys are kept based on the FLAG entry and a survey hierarchy set in SD18 (APOGEE, GALAH, GES, RAVE-ON, RAVE, LAMOST, and SEGUE). For example, if APOGEE and RAVE contain a duplicate entry with proper input and output data (FLAG = 0) then, regardless of the relative errors, the APOGEE entry is kept and labelled as DUPLICATED = 0. Lastly, the BEST flag in

SD18 encompasses the two prior flags and is 1 when FLAG = 0, DUPLICATED = 0 and the star has a valid cross-match in *Gaia* DR2.

For this work, we choose a sub-sample of stars from the SD18 data set by applying a number of quality cuts. In particular, we choose only those stars for which distance (d) is less than 10 kpc, $1/\text{parallax} < 10$ kpc, BEST = 1, and the ratio of parallax to error in parallax, $\pi/\sigma_\pi > 3$. Since the parallax/distance bias is more significant for very young stars (for details see Sanders & Das 2018), we discard all the stars with ages less than 1 Gyr. Our selected sub-sample has a total of 3147 069 stars.

5.2 Comparison of SD18 data set with a bias-controlled data set

Before we proceed to the calculation of breathing motions for the Milky Way using our selected sub-sample from the SD18 data set, first we compare our sub-sample with a bias-controlled data set for the Milky Way. We choose the Schönrich et al. (2019) data set (hereafter SME19) who corrected for the bias in the *Gaia* parallax, and derived the distances of all stars in the RVS (radial velocity spectrograph) sample of the *Gaia* DR2 via a Bayesian framework (for details, see Schönrich et al. 2019).

Since the SME19 data set does not include stellar ages, we do not include any age cuts on the SD18 data set for this comparison. In both data sets, we select stars with distance < 10 kpc and with the ratio of parallax to parallax error, $\pi/\sigma_\pi > 3$. Furthermore, to reduce the contamination from both very young stars and halo stars (at larger heights), we further constrain the sample based on the radial action (J_R) of the individual stars. The SD18 data set includes the radial action of individual stars; for the SME19 data set we calculate these ourselves. To calculate J_R , we use the Stäckel fudge method (Binney

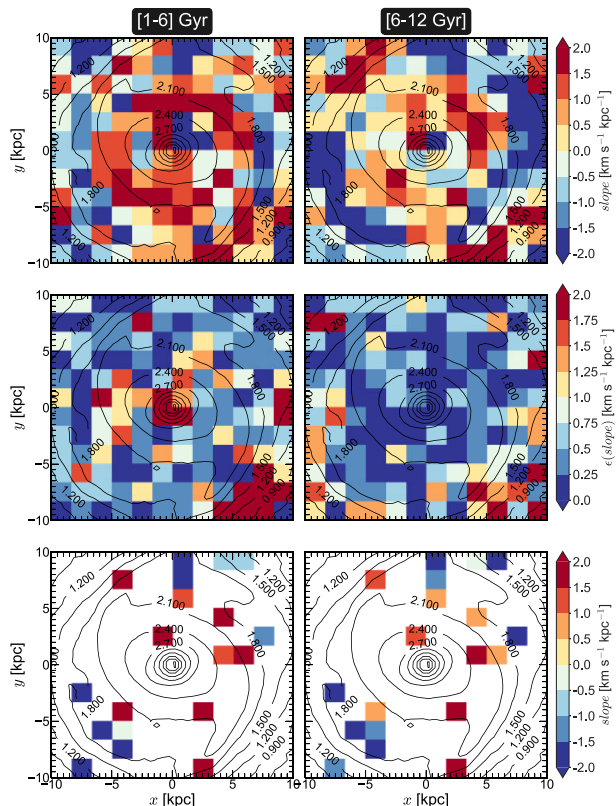


Figure 9. Star-forming model: distributions of the best-fitting slope (top panels) and the associated errors (middle panels) in the (x, y) plane for two populations with different ages are shown. Bottom panels show the same distribution of the best-fitting slope, but only for those bins where a simultaneous quality cut of $\text{slope}/\epsilon(\text{slope}) > 3$, for both the stellar populations, is met. Black solid lines denote the contours of constant density.

2012; Sanders & Binney 2016) from the AGAMA software library (Vasiliev 2019) in the potential of McMillan (2017). The Stäckel fudge function in AGAMA takes the Galactocentric 6D coordinates of the SME19 data set and produces the full action-angle coordinates, including the J_R values.

We then compute the cumulative distribution function (CDF) of the radial action, $F(J_R)$, of the SME19 data set. The very young stars are, in general, on almost circular orbits and thus should have small values of J_R . On the other hand, the halo stars at larger heights should have large values of J_R . Therefore, we choose stars from the SME19 data set, while rejecting the top and bottom 10 percent of stars in $F(J_R)$. We denote this rejection process as $\Delta F(J_R) = 0.1$. To make a uniform comparison, we apply the same radial action cut [$\Delta F(J_R) = 0.1$] on the SD18 data set using the same values of upper and lower J_R as in the SME19 data set. Fig. 11 shows the corresponding distributions of stars in the (x, y) plane and in the (R, z) plane, as well the distribution of mean vertical velocity, $\langle v_z \rangle$, in the (R, z) plane for both the SD18 and SME19 data sets. We split our selected samples of stars from both the data sets, into positive and negative Galactocentric azimuthal angles, and recalculate the distribution of $\langle v_z \rangle$, which is shown in Fig. 12. The distribution of the mean vertical velocity ($\langle v_z \rangle$) in the (R, z) plane, calculated for both the SD18 and SME19 data sets appear to match quite well. Further, we checked that a more conservative cut on J_R , e.g. $\Delta F(J_R) = 0.15$, does not alter the main findings presented here and in the subsequent sections.

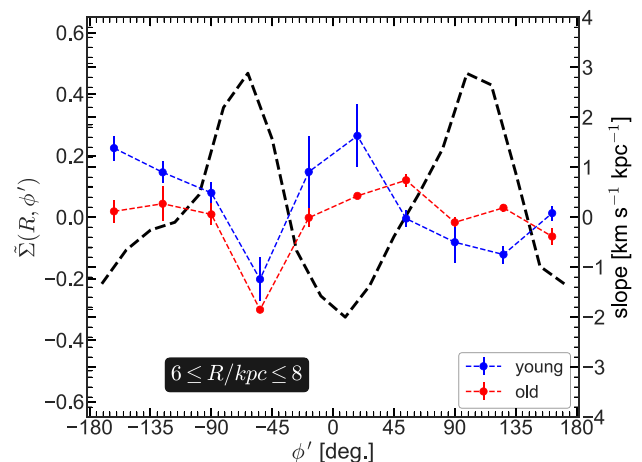


Figure 10. Star-forming model: variation of the slopes for the young (blue circles) and the old stellar populations (red circles) are shown as a function of the rotated azimuthal angle (ϕ'). Only the stellar particles in the radial range 6–8 kpc are chosen here. The black dashed line denotes the azimuthal variation of the total residual surface density $[\tilde{\Sigma}(R, \phi')]$, calculated in this chosen radial extent. The particles have first been binned in 1 kpc annuli and then azimuthally rotated relative to each other so that the minimum in $[\tilde{\Sigma}(R, \phi)]$ in each annulus is coincident. Only then are the vertical slopes computed, amounting to stacking different radial ranges while unwinding the spiral. The sense of rotation is towards increasing ϕ .

We calculate the distributions of the bending velocity (V_{bend}) in three vertical layers: $|z| = [0, 0.4]$ kpc, $|z| = [0.4, 0.8]$ kpc, and $|z| = [0.8, 1.2]$ kpc (using equation 5) for both the SD18 and SME19 data sets. The resulting bending velocities calculated from the two data sets are broadly in agreement; these results are not shown here but are evident in Fig. 11. We also compare the breathing motions, V_{breath} , between the two data sets. Because the breathing motions are significantly smaller than the bending motions, we start looking for the signature of the breathing motions in the Solar Neighbourhood in the following way. First, we calculate the distribution of the best-fitting slope in the (x, y) plane around the Solar Neighbourhood, by fitting a straight line to the variation of $\langle v_z \rangle$ with z , in a similar way as done for the model. The corresponding distribution of the best-fitting slope, and the associated error [$\epsilon(\text{slope})$] in the (x, y) plane, for both the SD18 and the SME19 data sets are shown in Fig. 13 (see top and middle panels). The bottom panels of Fig. 13 show the bins for which a quality cut of $\text{slope}/\epsilon(\text{slope}) > 3$ is met *simultaneously* for these data sets. We find that the location where we get the most consistent signature of the V_{breath} is $x \in [7.6, 8.1]$ kpc and $y \in [0.9, 1.4]$ kpc. In this bin, the best-fitting slope values, obtained from the two data sets agree with each other (within their error bars). Moreover, we find that this bin shows the largest value for the best-fitting slope, i.e. the signature of the largest V_{breath} . The presence of this ‘consistent’ signature of V_{breath} is not seen for any other bins shown in the bottom panels of Fig. 13. Hence, we will only consider this spatial bin for the subsequent study of the age dependence of the V_{breath} in the Solar Neighbourhood. The corresponding slope of $\langle v_z \rangle$ versus z at the same spatial location using the two data sets are shown in Fig. 14. We note that there is an offset between the $\langle v_z \rangle$ values calculated from both the data sets, especially near the disc mid-plane; this could possibly be due to the different zero-points used in the data sets. The values of the slopes obtained from the two data sets match quite well within their error bars. In both cases, the $\text{slope}/\epsilon(\text{slope})$ is found to be greater than 3 so that in both data sets the detection of V_{breath} is significant.

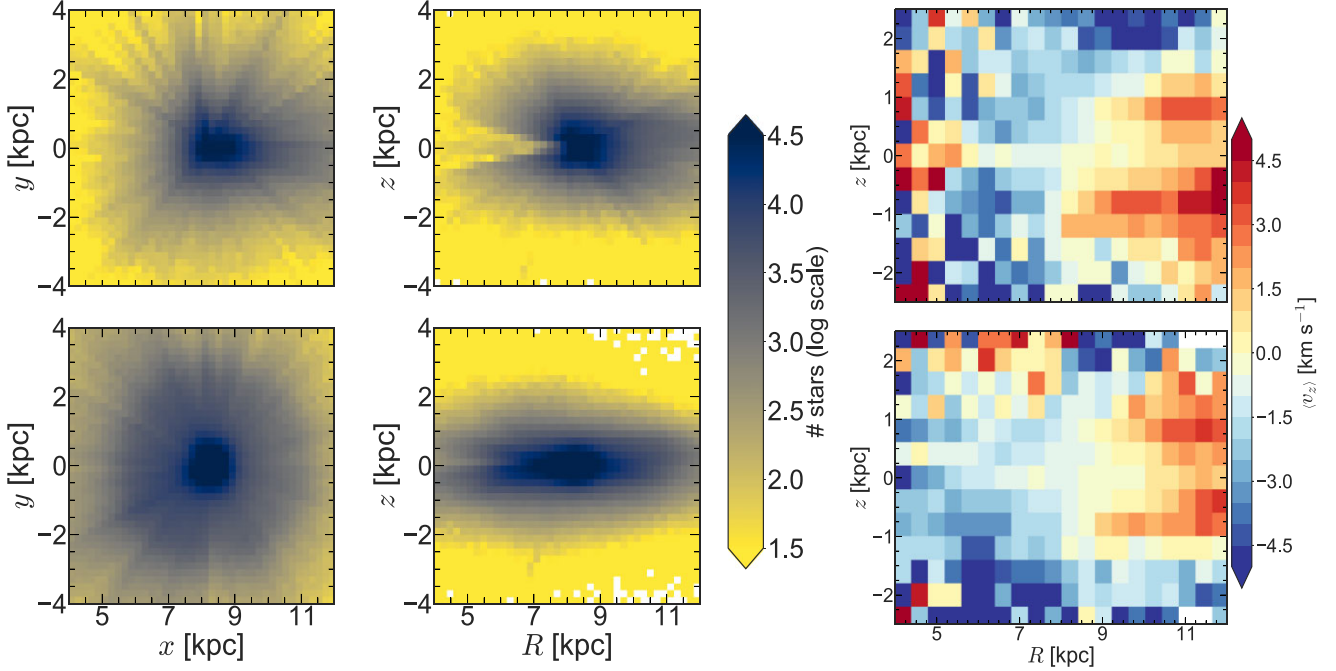


Figure 11. Comparison between the SD18 and SME19 data sets: density distribution of stars in the (x, y) -plane (left column) and in the (R, z) -plane (middle column), and the mean vertical velocity in the (R, z) -plane (right column) are shown for the SD18 (top panels) and the SME19 (bottom panels) data sets. For the details of the quality cut applied to both data sets, see Section 5.2.

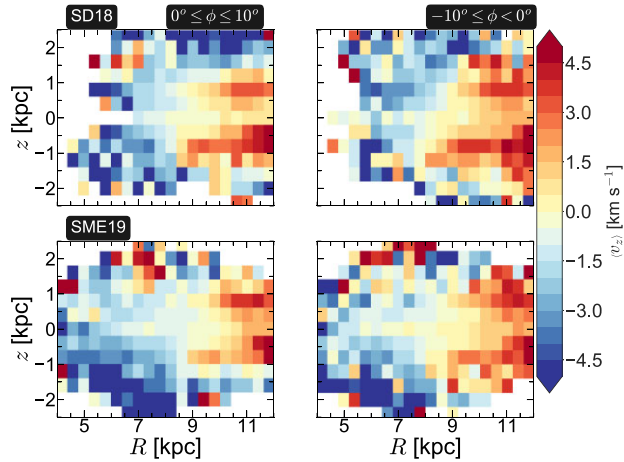


Figure 12. Comparison between the SD18 and SME19 data sets: distributions of the mean vertical velocity, $\langle v_z \rangle$, are shown for our selected samples of stars when they are split into positive and negative azimuthal angles. The top panels show results for the SD18 data set while the bottom panels show those for the SME19 data set. The extent of azimuthal range is provided at the top of each column.

5.3 Vertical motions for different ages in the Milky Way

We now study the non-zero mean vertical motion, and the associated breathing motions, as a function of the stellar ages in the same volume as where we found the largest breathing motions. We can do this only using the SD18 data set.

Fig. 15 presents how the selected sample of stars from the SD18 data set is divided into three equal-sized populations by cutting

the CDF, $F(\text{age})$, of the stellar ages. The resulting sub-populations are: young $[0 \leq F(\text{age}) < 1/3]$, intermediate $[1/3 \leq F(\text{age}) < 2/3]$, and old $[2/3 \leq F(\text{age}) < 1]$ stellar populations, or, in terms of the stellar ages, $[1, 4.85]$ Gyr (young), $[4.85, 7.1]$ Gyr (intermediate), and $[7.1, 12.6]$ Gyr (old). Fig. 16 shows the distributions of these stellar populations in the Galactocentric (x, y) plane, as well as in the (R, z) plane.

Fig. 17 shows the mean vertical velocity distribution in the (R, z) plane for the three stellar populations. Non-zero mean vertical motion is seen in all three stellar populations. Stars on either side of the mid-plane are seen to move in opposite directions, i.e. with a non-zero V_{breath} at larger heights, while at moderate heights all populations exhibit bending motions (i.e. stars on either side of the mid-plane are moving in the same direction).

To measure the variation of the breathing velocity as a function of stellar age, we consider the same spatial location as before, i.e. $x \in [7.6, 8.1]$ kpc and $y \in [0.9, 1.4]$ kpc, where a strong (and consistent) breathing motion is present without an age cut. The vertical distributions and slopes of the three populations in this spatial bin are shown in Fig 18. There is an offset in the zero-point among the different age groups, similar to what is seen in the model. We find a similar explanation as in the model, i.e. a difference in bending motion amplitudes in different age groups is giving rise to this offset. The presence of a non-zero slope in all three populations indicates that all of them take part in the breathing motion. The values of the best-fitting slope vary with stellar age: the (absolute) value of the best-fitting slope is largest for the intermediate stars ($2.27 \pm 0.25 \text{ km s}^{-1} \text{ kpc}^{-1}$) and smallest for the old stars ($1.02 \pm 0.19 \text{ km s}^{-1} \text{ kpc}^{-1}$) whereas the value of the best-fitting slope for the young star falls in between ($1.61 \pm 0.06 \text{ km s}^{-1} \text{ kpc}^{-1}$). We check whether the stars at the larger distance from the disc mid-plane bias the fits by ignoring the two points at the largest height

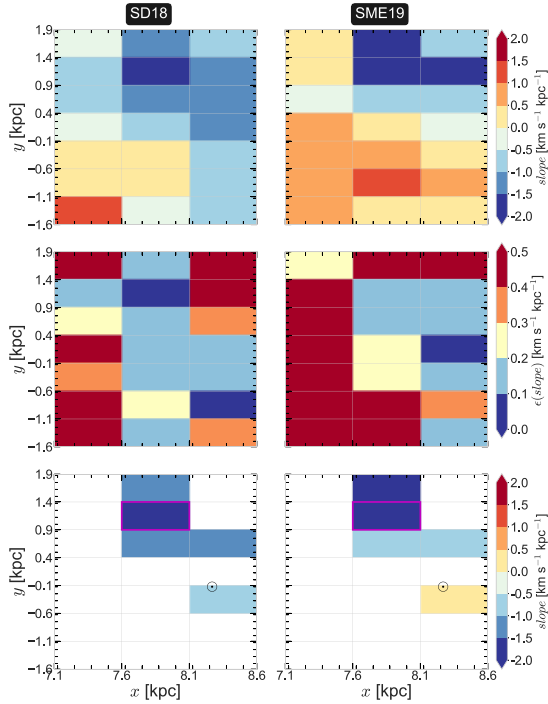


Figure 13. Comparison of the breathing motions in the SD18 and SME19 data sets: distributions of the best-fitting slope (top panels) and the associated errors (middle panels) in the (x, y) plane are shown for the SD18 (left-hand panels) and the SME19 (right-hand panels) data sets. Bottom panels show the same distribution of the best-fitting slope, but only for those bins where a simultaneous quality cut of $\text{slope}/\epsilon(\text{slope}) > 3$, for both the data sets, is met separately. The magenta rectangles in the bottom panels denote the chosen spatial bin where we study the age variation of the breathing motions (see text for details). The Sun's position is also indicated in the bottom panels.

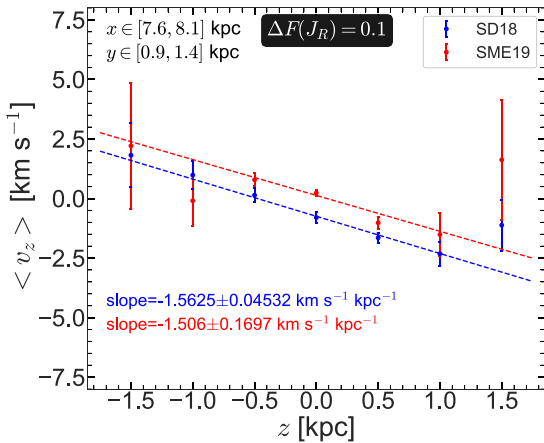


Figure 14. Comparison of the breathing motions in the SD18 and SME19 data sets. The best-fitting straight line fits to the $\langle v_z \rangle$ versus z data are shown for both data sets. The spatial location where the slopes are measured, indicated at the top of the figure, is chosen to maximize the slope in the SD18 data set. The corresponding best-fitting slope and the associated error are indicated in the top left label. Details of the quality cuts applied to both data sets are given in Section 5.2. The error bars on $\langle v_z \rangle$ are the standard errors ($\sigma_{v_z}/\sqrt{N}$) on the mean vertical velocity.

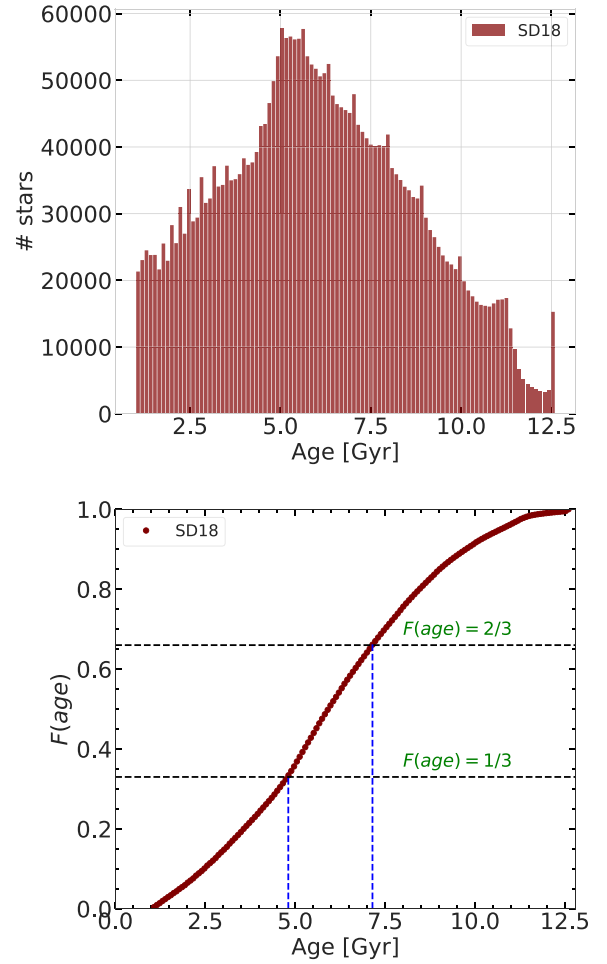


Figure 15. SD18 sample: histogram of stellar ages (top panel) and the corresponding CDF, $F(\text{age})$ of stellar ages (bottom panel) of our selected sample from the SD18 data set. The black dashed lines in the bottom panel indicate the cuts applied on the CDF whereas the vertical blue dashed lines indicate the corresponding stellar ages.

($|z| = 1.5$ kpc). The resulting best-fitting slopes (-1.6 ± 0.1 km s $^{-1}$ kpc $^{-1}$ for young, -2.26 ± 0.6 km s $^{-1}$ kpc $^{-1}$ for intermediate, and -1.5 ± 0.04 km s $^{-1}$ kpc $^{-1}$ for old stellar populations) match, within the error bars, with the slopes obtained including the largest heights, indicating that the large $|z|$ points are not biasing the fits.

To test the robustness of the best-fitting slope to systematic biases in the *Gaia* distances, we lower the distances of all the stars in our selected sub-sample by 10 per cent to simulate a parallax bias. After repeating the analysis of fitting a straight line to the variation of $\langle v_z \rangle$ with height, the resulting change in the best-fitting slope is negligible (less than 5 per cent) for all three age populations. This emphasizes that the breathing motions are not sensitive to systematic errors of this type and accentuating the robustness of the detection of breathing motions in the Solar Neighbourhood.

So far, we have estimated the statistical errors on $\langle v_z \rangle$ and also checked for the systematic errors on the best-fitting slope. However, the uncertainties associated with the parallax and the proper motions, the quantities used to derive vertical motion of stars, also result in uncertainties in the calculation of the $\langle v_z \rangle$ and the corresponding best-fitting slope. To estimate this error on the best-fitting slope, we perform a Monte Carlo analysis. First, we calculate the slope and the

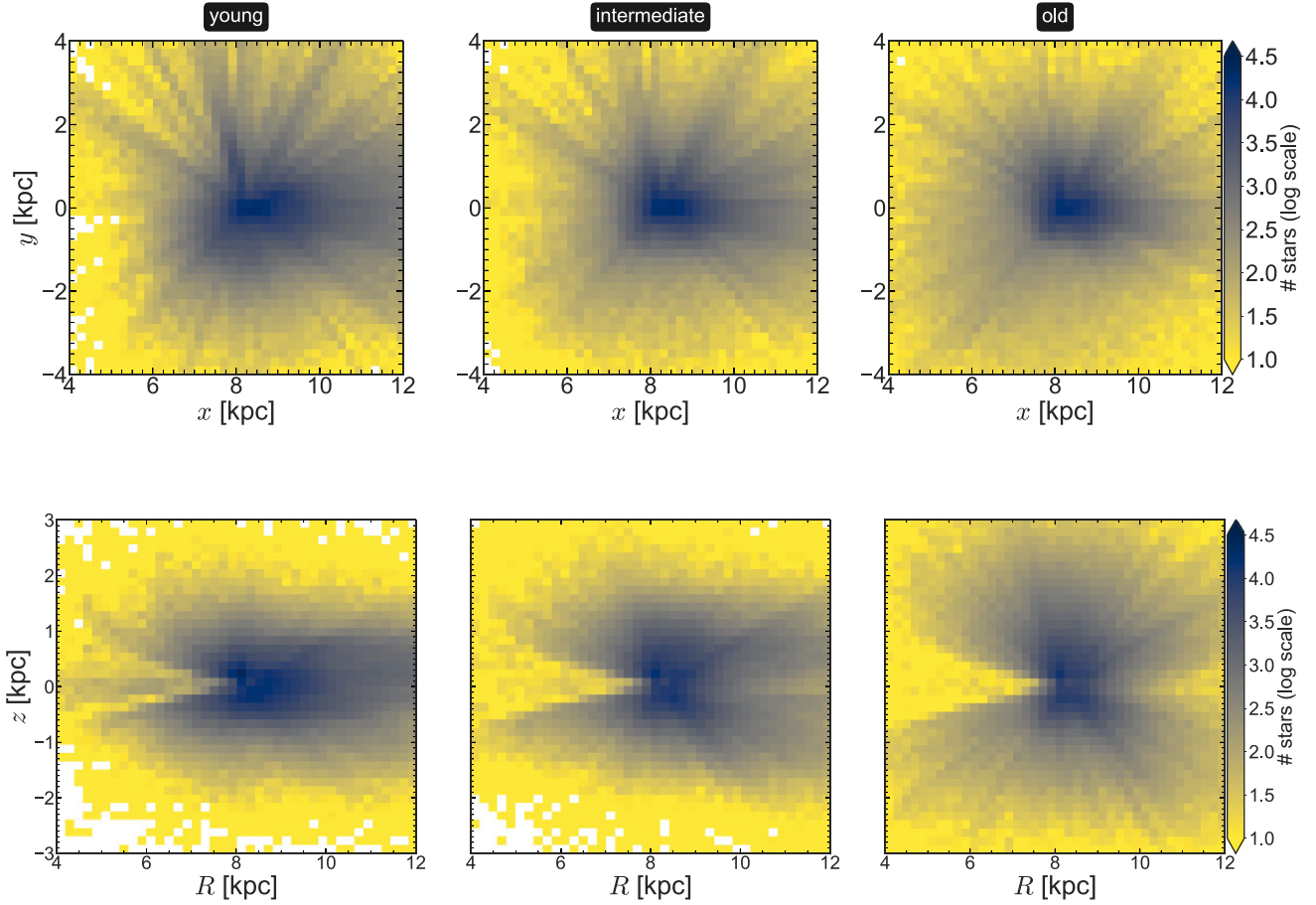


Figure 16. SD18 sample: distribution of our selected sample from SD18 is shown in the (x, y) plane (top panels) and in the (R, z) plane (bottom panels) for three stellar populations: young (left-hand panel), intermediate (middle panel), and old (right-hand panel) stellar populations (for details, see Section 5.3). An additional cut on the radial action (J_R) is applied; for details, see text in Section 5.2.

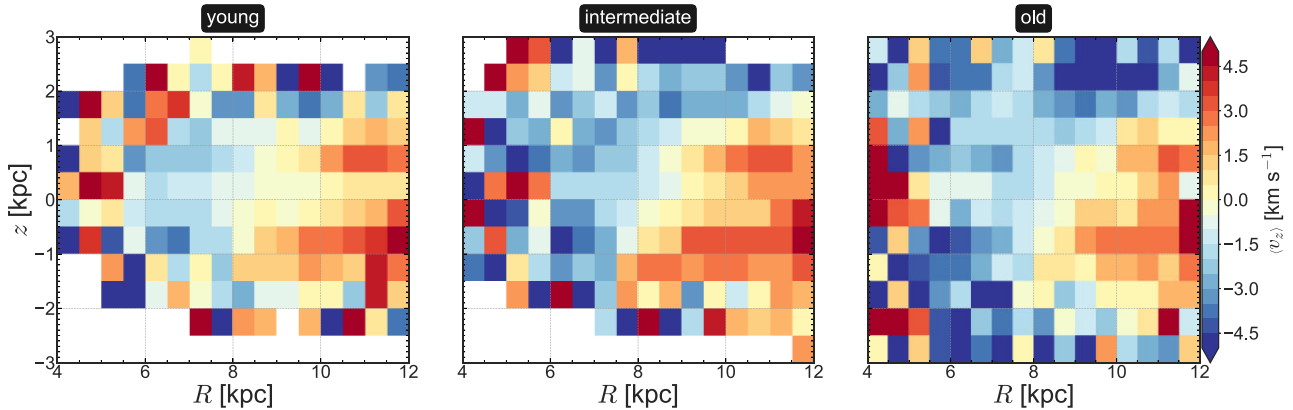


Figure 17. SD18 sample: mean vertical velocity, $\langle v_z \rangle$, for different stellar populations in the (R, z) plane. From left to right: stars with ages 1–4.85 Gyr (young), 4.85–7.1 Gyr (intermediate), and 7.1–12.6 Gyr (old), are shown (for details, see Section 5.3). An additional cut on the radial action (J_R) is applied; for details, see text in Section 5.2.

associated error of the best-fitting straight line using the 6D position–velocity of stars in our full selected sub-sample. Since there is a robust signal of breathing motions only in our selected spatial location in the Solar Neighbourhood, we restrict our Monte Carlo analysis for that spatial location only. Then, using the errors in parallax and the

proper motion, we add random Gaussian errors in the parallax and the proper motion of each star in our sub-sample, and recalculate the 6D position–velocity of stars in our selected sub-sample using GALPY (Bovy 2015). Once we recalculate the 6D position–velocity of stars, we repeat the exercise of fitting a straight line to the variation

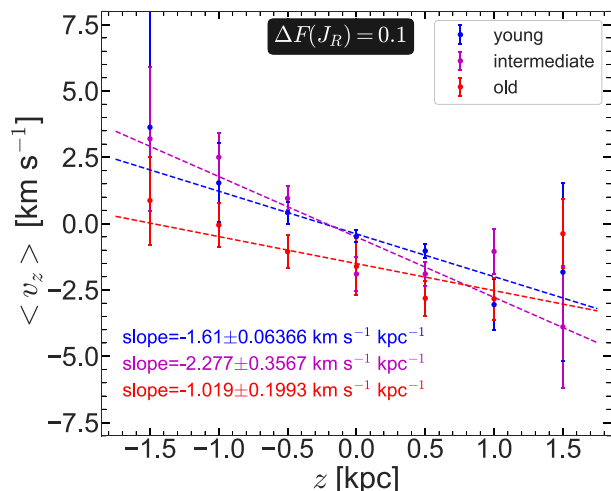


Figure 18. SD18 sample: The best-fitting straight lines to the $\langle v_z \rangle$ versus z data are shown for three populations with different ages, calculated at the spatial location $x \in [7.6, 8.1]$ kpc and $y \in [0.9, 1.4]$ kpc. The corresponding best-fitting slope and the associated error are mentioned in the figure. The intermediate stellar population has the largest slope while the old stellar population shows the smallest slope. The error bars on $\langle v_z \rangle$ are the standard errors ($\sigma_{v_z}/\sqrt{N}$) on the mean vertical velocity.

of $\langle v_z \rangle$ with height, z , for all three stellar populations with different ages. This process is repeated 1000 times. The resulting distributions of best-fitting slopes, calculated for that particular spatial bin, for all three stellar populations are shown in Fig. 19. The estimated errors in the best-fitting slope measurements, obtained from the Monte Carlo analysis are ~ 14 per cent for the young population, ~ 16 per cent for the intermediate population, and ~ 17 per cent for the old population.

6 SIGNATURES OF SPIRAL DRIVEN BREATHING MOTIONS

Using a star-forming simulation, we have shown that a spiral density wave drives vertical breathing motions. The increasing amplitude of the breathing velocity with height above the mid-plane, previously reported by Debattista (2014), is consistent with the trend seen in the Milky Way (Gaia Collaboration 2018a). This behaviour is

an important clue to the source of breathing motions: it must be something local to affect stars near the mid-plane differently from those further up. The reason why spirals have this effect is easy to understand: on vertical scales small compared to the disc scale length, the vertical gravitational field can be approximated, using Gauss's law, as a uniform field from an infinite sheet of mass corresponding to the mass enclosed within a cylinder of the same height and small radius. Both the simulation of Debattista (2014) and the one considered here show that the spiral density wave extends to large heights (see Fig. 3). If the density of the disc is decomposed into an axisymmetric part and a spiral perturbation part:

$$\rho(R, \phi, z) = \rho_0(R, z) + \rho_s(R, \phi, z), \quad (6)$$

then the vertical velocities are in equilibrium with the axisymmetric part of the potential and $\rho_0(R, z)$ gives rise to $V_{\text{breath}} = 0$. Since the azimuthally varying part, ρ_s , is due to spirals, it extends to large height. The enclosed mass therefore increases with height as the vertical integral of $\rho_s(R, \phi, z)$. Therefore the gravitational field increases with height, as we see in the simulations, and in the Milky Way (compare Fig. 6 here and fig. C.6 in Gaia Collaboration 2018b). It is noteworthy that this vertical increase in V_{breath} is absent in the model of Faure et al. (2014), who used an analytic spiral potential with a very short scale height of 100 pc, much shorter than that of the unperturbed density distribution. As a result, the gravitational field reaches nearly constant value quite close to the mid-plane in their model. Consequently, V_{breath} shows a near-linear increase from the disc mid-plane, and after $\sim 0.7z_0$, it starts to decline where z_0 is the scale height of the spiral perturbation. The vertical variation of V_{breath} in the *Gaia* data therefore indicates that the spiral structure in the Milky Way is vertically extended. The vertical increase of V_{breath} is different from the nearly constant V_{bend} as a function of height (compare both panels of fig. C.6 in Gaia Collaboration 2018b). This indicates that the Milky Way's bending motions are probably not excited by spirals.

Moreover, we examined the dependence of the breathing velocity on stellar age. The star-forming model exhibits a decreasing V_{breath} with increasing stellar age (see Fig. 10) indicating that increasing random motions decrease the vertical response of the stars. Fig. 18 shows that, using the SD18 data set, the intermediate stellar population indeed yields a larger (absolute) value of the best-fitting slope in the Solar Neighbourhood compared with the old stellar

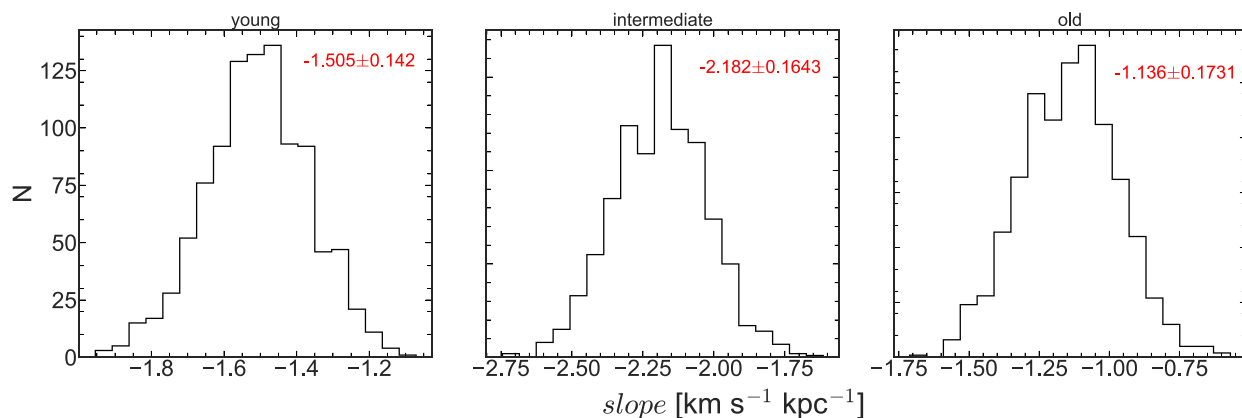


Figure 19. SD18 sample: distribution of the best-fitting slopes, calculated in the Monte Carlo analysis (for details, see Section 5.3) are shown for three stellar populations with different ages, at the spatial location $x \in [7.6, 8.1]$ kpc and $y \in [0.9, 1.4]$ kpc. The median values and the associated errors are indicated in each sub-panel.

population, while the young stellar population has a slope that falls in between; however, the difference between the young and intermediate age populations (1.8σ) is not statistically significant in this data set. In other words, the breathing motion is stronger for the younger stellar populations (i.e. young and intermediate populations considered together) than that for the old stellar population in the Solar Neighbourhood.

For these reasons, we therefore argue that the breathing motions seen in the Solar Neighbourhood are very likely to be driven by the spiral density wave structure. Widrow et al. (2014) proposed a scenario where a satellite galaxy, while plunging into the disc, can generate both bending and breathing motions, depending on the value of the satellite's velocity in the direction perpendicular to the disc. Such interactions also excite a strong spiral response within the disc that are probably partly responsible for the breathing motions seen in the simulations of Widrow et al. (2014). The vertical variation of the amplitude of the breathing modes excited by satellites is also unclear.

7 FUTURE PROSPECTS AND SUMMARY

Our requirement of stellar ages limits the number of stars available to us by almost a factor of 2 when compared with the *Gaia* MAIN SAMPLE (which contains 6376 803 stars; Gaia Collaboration 2018b). In addition, the azimuthal coverage in *Gaia* DR2 is not large ($\Delta\phi \sim 30^\circ$). This, in turn, prevents us from exploring the change from expanding to compressing breathing motions as a function of azimuthal angle, which would help constrain the extent of the spiral and its corotation radius (Debattista 2014; Faure et al. 2014). *Gaia* DR3 holds the potential to permit us to investigate the variation of breathing motions as a function of azimuth, in order to constrain the Milky Way's spiral structure better.

Our main findings are as follows:

(i) The spirals in the self-consistent star-forming model drive coherent vertical breathing motions, with amplitude that increases with the height from the disc mid-plane. We argue that this vertical variation is a consequence of the vertically extended spiral structure. Since *Gaia* DR2 also exhibits this increasing breathing amplitude with height, it implies that the Milky Way's spirals are also vertically extended.

(ii) Prominent breathing motions are present in stellar populations of all ages in the model, with the strongest motions in the young stellar population, and weakest in the old one.

(iii) Using *Gaia* DR2 with complementary age data from Sanders & Das (2018), we showed that the location near the Solar Neighbourhood with the smallest biases shows significant, strong breathing motions for all ages. In addition, the amplitude of the breathing motions varies with stellar age, with the oldest population displaying the weakest breathing amplitude (when compared with the intermediate and the young population, at $2.8-3.1\sigma$). The difference between the young and intermediate age populations is smaller but less statistically significant (1.8σ). This behaviour is similar to the trend shown in the simulation model and is consistent with spirals being the driving mechanism of breathing modes.

The resemblance in the height and age variation of the amplitude of breathing velocity, in the star-forming model and in the location in the Milky Way showing prominent breathing motions from the *Gaia* DR2 data set indicates that the observed breathing motions in the Milky Way are likely to be excited by spiral density waves. Due to the limited coverage and uncertainties in the *Gaia* parallax, we have restricted to the Solar Neighbourhood our search for breathing

motions, but with the upcoming *Gaia* DR3, there will be a larger number of stars allowing us to improve the statistics in the Solar Neighbourhood.

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DATA AVAILABILITY

The simulation data underlying this article will be shared on reasonable request to VPD (vpdebattista@gmail.com). The Sanders and Das data set is publicly available from this URL.

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